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An evolutionary system of mineralogy, part II: Interstellar and solar nebula primary condensation mineralogy (> 4.565 Ga)

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ABSTRACT

The evolutionary system of mineralogy relies on varied physical and chemical attributes, including trace elements, isotopes, solid and fluid inclusions, and other information-rich characteristics, to understand processes of mineral formation and to place natural condensed phases in the deep-time context of planetary evolution. Part I of this system reviewed the earliest refractory phases that condense at T > 1000 K within the turbulent expanding and cooling atmospheres of highly evolved stars. Part II considers the subsequent formation of primary crystalline and amorphous phases by condensation in three distinct mineral-forming environments, each of which increased mineralogical diversity and distribution prior to the accretion of planetesimals > 4.5 billion years ago:

1) *Interstellar molecular solids*: Varied crystalline and amorphous molecular solids containing primarily H, C, O, and N are observed to condense in cold, dense molecular clouds in the interstellar medium ($10 < T < 20$ K; $P < 10^{-13}$ atm). With the possible exception of some nano-scale organic condensates preserved in carbonaceous meteorites, the existence of

these phases is documented primarily by telescopic observations of absorption and emission spectra of interstellar molecules in radio, microwave, or infrared wavelengths.

2) *Nebular and circumstellar ice*: Evidence from infrared observations and laboratory experiments suggest that cubic H_2O ("cubic ice") condenses as thin crystalline mantles on oxide and silicate dust grains in cool, distant nebular and circumstellar regions where T ~100 K.

3) *Primary condensed phases of the inner solar nebula*: The earliest phase of nebular mineralogy saw the formation of primary refractory minerals that solidified through high-temperature condensation ($1100 < T < 1800$ K; $10^{-6} < P < 10^{-2}$ atm) in the solar nebula more than 4.565 billion years ago. These earliest mineral phases originating in our solar system formed prior to the accretion of planetesimals and are preserved in calcium-aluminum-rich inclusions, ultra-refractory inclusions, and amoeboid olivine aggregates.

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INTRODUCTION

The incremental, episodic emergence of mineral diversity and distribution through more than 13 billion years of cosmic evolution provides the basis for an "evolutionary system" of mineral classification—one that emphasizes the formation of solid phases by a progression of physical, chemical, and ultimately biological processes (Hazen et al. 2008; Hazen and Ferry 2010; Hazen 2019; Hazen and Morrison 2020). This system amplifies the official classification protocols of the International Mineralogical Association's Commission on New Minerals, Nomenclature and Classification (IMA, CNMNC; e.g., Burke 2006; Mills et al. 2009; Schertl et al. 2018), which defines each mineral "species" on the basis of its unique combination of end-member composition and idealized crystal structure. More than 5500 approved mineral species are now recognized by the IMA system (rruff.info/ima; accessed 27 March 2020).

By design, the IMA classification system is predicated on identifying the minimum amount of information (as measured in bits; e.g., Krivovichev 2012, 2013) required to distinguish one species from another. Consequently, IMA procedures cannot capture the information-rich complexity of natural mineral specimens—their trace and minor elements, fractionated isotopes, structural defects, varied magnetic and electrical properties, external morphologies, solid and fluid inclusions, spectral features, petrologic environment, ages of both formation and subsequent diagenetic episodes, and myriad other attributes that have the potential to tell the story of each individual sample's origin and alteration via interactions with a succession of environments through deep time. We conclude that IMA protocols are insufficient to classify minerals in their evolutionary contexts.

Accordingly, we propose an "evolutionary system of mineralogy" that amplifies and modifies the IMA scheme in three ways, each of which is informed by those information-rich aspects of

natural mineral specimens—attributes that are the essence of historical science discovery in the "messy, uncontrollable world of nature" (Cleland 2013; see also Cleland 2011). We split some IMA species into two or more "natural kinds"—subdivisions that recognize fundamentally different idiosyncratic combinations of attributes that arise from distinct paragenetic modes. Thus, for example, we view isotopically anomalous nanoscale diamond condensed from a high-temperature, low-pressure carbon-rich vapor in the expanding atmosphere of an exploding star as intrinsically different from macroscopic Type I "gem" diamond crystallized in high-temperature, high-pressure, carbon-saturated aqueous fluids in Earth's mantle, which in turn differs from diamond formed by the impact of an asteroid on near-surface carbonaceous material (Hazen 2019). Such splitting of IMA species into multiple "natural kinds" is appropriate for many of Earth's commonest mineral species, including calcite, hydroxylapatite, pyrite, and quartz, all of which have both abiotic and biotic paragenetic modes.

In other instances, we propose lumping two or more IMA species into a single "natural kind." Notable examples of species that are lumped according to their evolutionary contexts occur in chemically diverse structural groups of rock-forming silicates, including amphibole, mica, pyroxene, and tourmaline group minerals, for which a given specimen formed in a single geological setting and in one continuous phase domain may display a range of zoning and solid solution that overlaps the compositional ranges of two or more end-member species as defined by IMA protocols. Other examples of IMA species that we lump into a single natural kind include groups of isostructural rare earth element minerals, some Mg-Fe oxides and silicates, and natural metal alloys, for which small variations in the ratios of crystal chemically similar elements that occur within one paragenetic environment may require multiple IMA end-member mineral species.

Thirdly, the evolutionary system catalogues a variety of non-crystalline or aperiodic condensed phases, including glasses (e.g., stellar amorphous carbon, impact maskelynite, volcanic obsidian) and mixed-phase nano-materials (bauxite, coal, and limonite, for example)—materials not usually considered in the current IMA scheme (e.g., Hazen et al. 2013, Table 3), though many of these phases were included in the revolutionary third edition of James Dwight Dana's *System of Mineralogy* (Dana 1850), from which the modern IMA formalisms have evolved.

Meteorite minerals

An important illustration of the demonstrable benefits and potential pitfalls of the IMA system is provided by the diversity, distribution, and modes of formation of meteorite minerals, as reviewed by Rubin and Ma (2017, 2020), who tabulate more than 400 types of minerals known from meteorites, the majority of which (including dozens of intriguing micro- or nano-scale phases discovered by Dr. Chi Ma) have been approved as official IMA species. However, more than 50 of the listed meteorite phases are not officially recognized by the IMA (see rruff.info/ima; accessed 27 March 2020), because they fail to meet IMA criteria for legitimate species owing to a range of reasons.

At least a dozen minerals in Rubin and Ma's list, including native Mo, Nb, and Re; carbides of Fe, Mo, and Zr; Ca and Nb oxides, and other primarily micron-scale phases, may eventually receive approval but have not yet been accepted by the IMA. However, many of Rubin and Ma's listed phases do not meet IMA requirements for end-member compositional variants of natural crystalline phases. In some cases, they list phases that effectively split IMA species as a consequence of additional minor elements. For example, they record "Ti-rich" and "V-rich" varieties of magnetite (in addition to magnetite), "pleonaste" (a varietal name for Mg-Fe^{2+} oxide

spinel), "sodium phlogopite," "Ca-armalcolite," and "carbonate-fluorapatite." They also split graphite into two polytypes—graphite 2H and graphite 3R.

On the other hand, in some instances Rubin and Ma (2017, 2020) lump mineral species into broadly inclusive mineral groups—e.g., apatite, feldspar, mica, olivine, and orthopyroxene. They also implicitly lump two or more approved minerals with end-member compositions into one solid solution; for example, their list includes "magnesiowüstite," which is an unapproved name for the solid solution between periclase (MgO) and wüstite (FeO), both of which end members are also listed. Similarly, IMA-approved magnesite ($MgCO_3$) and siderite ($FeCO_3$) are listed along with "breunnerite," an obsolete, though useful, name for Mg-dominant rhombohedral carbonates in the magnesite-siderite solid solution. "Plagioclase" also appears in the tabulation along with the two end members of the albite-anorthite solid solution. Similarly, "biotite," "Au-dominated alloys," and "PGE-dominated alloys" are names for chemically complex phase regions that encompass multiple approved IMA species.

Finally, Rubin and Ma (2017, 2020) catalog several amorphous or composite nanophases that are excluded from the current IMA list (illite, martinsite, and maskelynite, for example). In these particulars, Rubin and Ma's tabulation of meteorite minerals mirrors the evolutionary system of mineralogy, as it underscores both the need for the exacting IMA nomenclature and the desirability of additional classification protocols to address complex natural condensed phases – varied split, lumped, and non-crystalline phases that are critical to describing the diverse modes of mineral formation in the natural world.

The potential benefits of the evolutionary system are further highlighted by Rubin and Ma (2017), who list 17 different processes by which the hundreds of documented meteorite minerals have formed (Table 1). They note, "Some meteoritic minerals form by only a single mechanism

(e.g., ringwoodite and ahrensite by high-pressure shock metamorphism of olivine); other minerals form by several mechanisms (e.g., olivine by condensation around red giant and AGB stars, condensation in the solar nebula, crystallization in [high-temperature nebular] melts, crystallization in chondrule melts, thermal metamorphism, crystallization from impact melts, condensation within impact plumes, crystallization in magmatic bodies on differentiated asteroids, annealing of amorphous material, and aqueous alteration)" (Rubin and Ma 2017, page 339). They emphasize that, "In this overview, no attempt has been made to describe every mode of formation of every meteoritic mineral. That monumental task would require a multi-volume book-length treatment."

We suggest that such varied modes of formation usually result in diagnostic sets of attributes that point to multiple natural kinds—the mineralogical key to understanding the evolution of planetary systems. Our multi-part endeavor, though far from the encyclopedic "monumental task" envisioned by Rubin and Ma for meteorite minerals, nevertheless will attempt to enumerate mineral natural kinds in the context of their paragenetic modes.

Table 1. Seventeen paragenetic modes of meteorite minerals, after Rubin and Ma (2017). Note that these modes of formation are not mutually exclusive. "Part" refers to proposed divisions of the evolutionary system of mineralogy, which will appear in multiple, ongoing contributions.

Mode #	Description	Part
1	Condensation in the gaseous envelopes of stars (stellar mineralogy)	I
2	Condensation in the solar nebula	II
3	Crystallization in CAI and AOA melts	II
4	Crystallization in chondrule melts	III
5	Exsolution during cooling of CAIs	II
6	Exsolution during cooling of chondrules	III
7	Annealing of amorphous material	I, II, & III
8	Thermal metamorphism and exsolution	V
9	Aqueous alteration, hydrothermal alteration, and metasomatism	V
10	Shock metamorphism	IV
11	Condensation within impact plumes	IV
12	Crystallization from melts in differentiated bodies	IV
13	Condensation from late-stage vapors in differentiated bodies	IV
14	Exsolution, inversion, and subsolidus redox effects in cooling igneous rocks	IV

15	Solar heating near perihelion	III,V
16	Atmospheric passage	V
17	Terrestrial weathering*	VI+

*Includes biological alteration, to be discussed in later contributions.

Pre-terrestrial stages of mineral formation

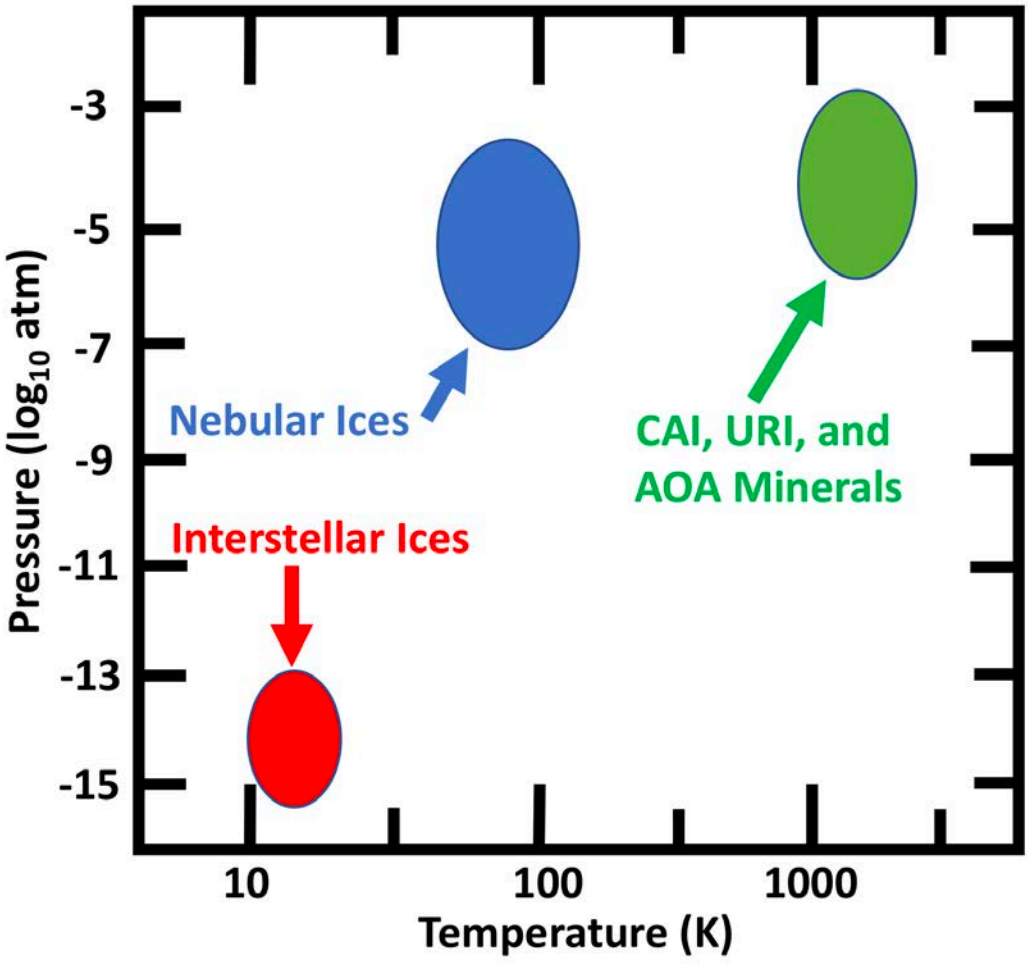
In Part I of this series, Hazen and Morrison (2020) described the earliest phase of mineral evolution commencing more than 13 billion years ago – the condensation of more than 40 kinds of extremely refractory nanoscale minerals, representing 22 IMA-approved mineral species, as well as two as yet unapproved crystalline phases and amorphous forms of carbon, alumina, and silicates (i.e., Mode #1 of Rubin and Ma's 17 modes of meteorite mineralization; Table 1). These varied phases were dispersed as micro- and nanoscale dust grains into the interstellar medium, along with hydrogen, helium, and other atomic and molecular species. A variety of highly evolved stars, including C- and O-rich asymptotic giant branch (AGB) stars, classical novae, and type II supernovae, seeded the interstellar environment with copious quantities of gas and mineral dust. Significant fractions of those raw materials eventually concentrated in subsequent generations of stars and their companion planets.

Here in Part II we examine the next two episodes of cosmic mineral evolution, which encompass the primary condensation, melt crystallization, and initial solid-state transformation on cooling (e.g., exsolution, reversible phase transitions, and element ordering) of a variety of crystalline and amorphous phases from the dust and gaseous remnants of stars. These processes include modes of formation #2, #3, and #5 in Table 1, as well as additional processes that lead to interstellar and circumstellar condensed phases that are observed through telescopic spectral observations but are not preserved in meteorites.

In particular, we consider a few kinds of condensed molecular phases, sometimes collectively known as "ices," though in this review the term ice refers exclusively to crystalline H_2O. Condensed molecular solids are observed by astronomical spectroscopy to form in the extremely low-temperature, high-vacuum conditions of "cool, dense" interstellar molecular clouds – the nurseries of new generations of stars. Perturbations to a molecular cloud can lead to a Jeans instability – a local density increase, followed by gravitationally-induced collapse of the molecular cloud into a solar nebula, and ultimately star system formation (Jeans 1902; Mizuno et al. 1994; Longair 2008). In the case of our solar system, more than 99.8 percent of the solar nebula's mass formed the Sun, commencing ~4.567 billion years ago, while the remaining dust and gas experienced a complex history of thermal processing and chemical mixing that led to the formation of numerous additional solid phases (Rubin and Ma 2017, 2020).

In this contribution we focus on presumed primary phases—those that are thought to have formed by condensation directly from a gas phase, by subsequent solidification from a melt at pressures less than 0.01 atm, and solid-state transformations (e.g., exsolution, annealing, or ordering), as observed in calcium-aluminum-rich inclusions (CAIs), amoeboid olivine aggregates (AOAs), and ultra-refractory inclusions (URIs) of chondrite meteorites (Figure 1). We also include condensation of cubic ice (H_2O), the low-temperature crystalline polymorph of water ice (e.g., Gaffney and Matson 1980), as < 0.05 micron-thick crystalline mantles on oxide and silicate dust particles in the cold, distant circumstellar regions of solar nebulae. It should be emphasized, however, that most primary condensed interstellar and solar nebular phases were either lost through sublimation or were eventually incorporated into larger bodies—asteroids, comets, and planetesimals that experienced subsequent alteration by large-scale differentiation; fluid-rock interactions; multiple thermal events, including heating by the decay of short-lived radioactive

isotopes, radiative heating, and conversion of gravitational potential energy; and high-energy impacts. Thus, a continuum exists between pristine "primary" and altered "secondary" mineral grains (e.g., Brearley and Jones 1998; MacPherson 2014; Rubin and Ma 2017). Primary mineral phases preserved in chondrules, which formed by melt solidification and subsequent solid-state reactions (Brearley and Jones 1998; Scott and Krot 2014; Connolly and Jones 2016), are the subject of Part III of this series, whereas the rich variety of presumed secondary minerals, as processed in the dynamic environments of growing planetesimals and preserved in the diverse meteorites that fall to Earth, will be the focus of Parts IV and V.



A

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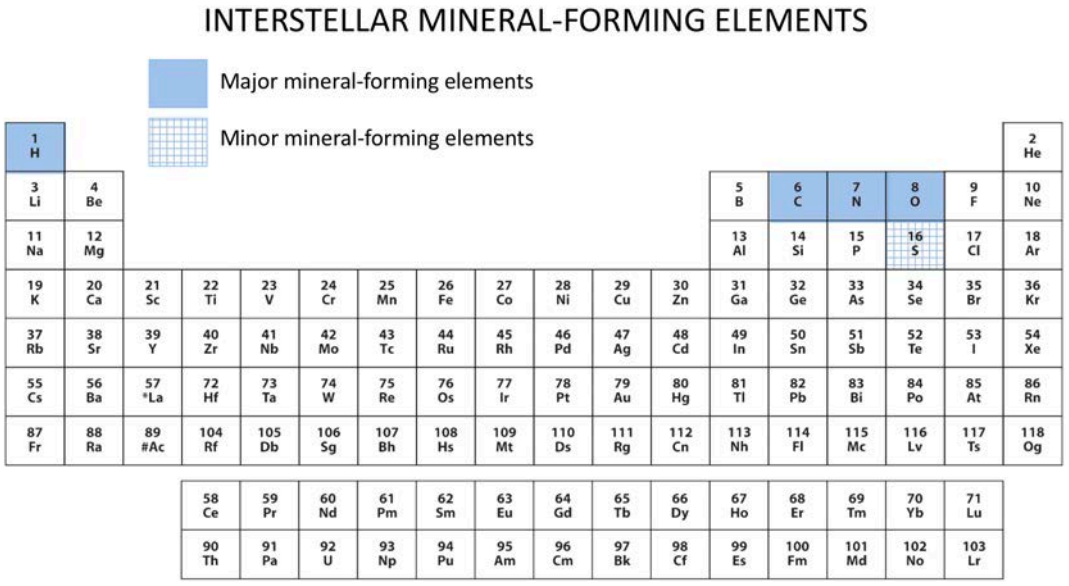
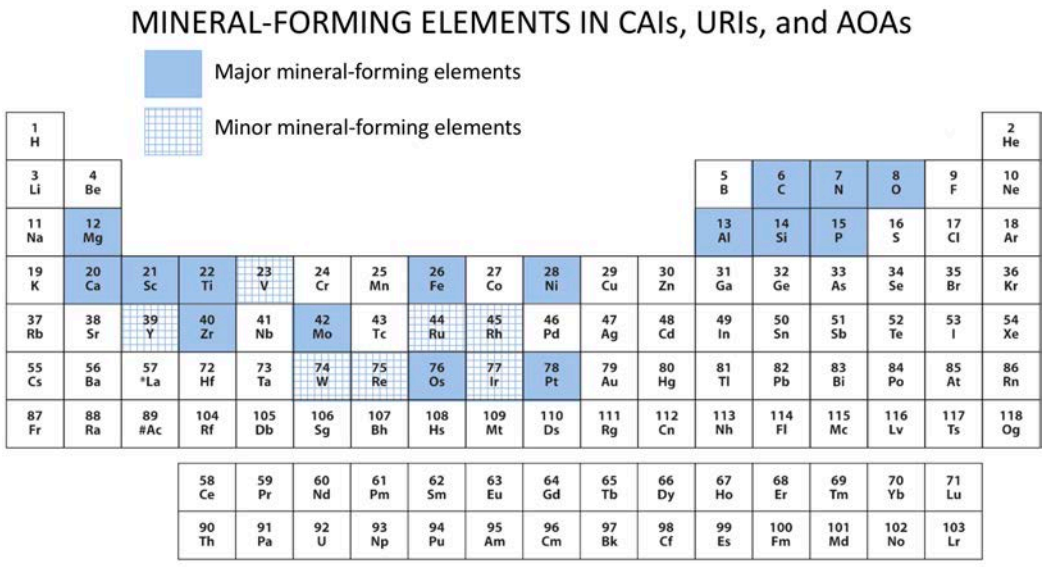


224 **B**



225 **C**

226 **Figure 1.** The temperature, pressure, and compositional characteristics of primary interstellar and
227 solar nebular condensates result in a distinctive second phase of the evolutionary system of
228 mineralogy. **A.** These minerals formed at a wide range of temperatures via low-pressure ($P < 0.01$
229 atm) condensation in interstellar and nebular environments. **B.** Interstellar minerals formed
230 primarily from C, H, N, O, and probably S – five of the most abundant elements in the cosmos. **C.**
231 Primary minerals in CAIs, URIs, and AOAs formed principally from 16 essential major elements,
232 with important additional contributions from 7 minor elements.

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THE MINERALOGY OF INTERSTELLAR MOLECULAR CLOUDS

Highly-evolved mineral-producing stars eject a significant fraction of their atmospheric dust and gas in energetic solar winds. Much of this ejected mass continues to be collected throughout the galaxy in a variety of molecular clouds, most of which are too warm and dispersed to form additional condensed phases (Greenberg 1991; Ferriere 2001). Only the so-called "cool, dense molecular clouds" have been the loci of a second significant pulse of cosmic condensation, albeit at significantly lower temperatures and pressures than their stellar precursors (Figure 2). These distinctive interstellar molecular clouds, typically composed of 99 percent gas, predominantly H in the form of H atoms and H_2 molecules (Wakelam et al. 2017) and atomic He, with important contributions by O-, C-, and N-bearing molecular species (Figure 1B), and ~1 weight percent (wt. %) mineral dust (primarily oxides, silicates, carbides, and carbon allotropes), are the nurseries of new star systems (e.g., Herbst 1995).

Figure 2. A 2015 Hubble Space Telescope image of a portion of the Eagle Nebula (NGC 6611 and IC 4703), dubbed "The Pillars of Creation," displays a star-forming region of a dense molecular cloud. The core regions of this structure are cooler areas, dense molecular clouds where molecular solids condense. [image credit: NASA, ESA and the Hubble Heritage Team (STScI/AURA)].

253

254 *The environments of cool, dense molecular clouds*

255 "Cool, dense molecular clouds" might seem misnamed. They are relatively cold ($10 < T < 20$
256 K) regions of the interstellar medium that are characterized by the gradual condensation of a
257 variety of molecular solids (Kouchi and Yamamoto 1995; Allamandola et al. 1999; Ferriere 2001;
258 Gibb et al. 2004). Furthermore, "dense" in this context equates to approximately 10^2 to a maximum
259 of 10^6 molecules per cubic centimeter—i.e., orders of magnitude lower pressure than the 10^{10}
260 molecules per cm^3 in a typical laboratory "high vacuum," though significantly greater than the
261 average galactic medium of fewer than one atom per cubic centimeter. Thus, the effective pressures
262 of condensation in cool, dense molecular clouds are thought to be less than 10^{-13} atm (Figure 1A).

263 Cool, dense molecular clouds are vast, sometimes exceeding 100 light years in diameter, and
264 they appear as dark irregular shapes in silhouette against the twinkling background of more distant
265 stars—a consequence of the cumulative effect of light-blocking dust (Ferriere 2001; Di Francesco
266 et al. 2006). Of note is the Taurus Molecular Cloud, which at a distance of ~430 light years is the
267 closest and among the best studied large star-forming regions (Luhman et al. 2010). In addition to
268 hundreds of relatively young stars, the Taurus Molecular Cloud hosts a diverse suite of molecular
269 species (Freeman and Millar 1983), some of which (though not all; e.g., Güston et al. 2019)
270 condense into molecular crystals and amorphous solids.

271 Evidence for such interstellar condensed phases comes primarily from infrared absorption
272 measurements, as light from more distant stars passes through a molecular cloud, as well as by IR
273 and radio emission spectroscopy. Astronomical observations are interpreted using spectroscopic
274 data from laboratory experiments on small molecules under cold vacuum (e.g., Allamandola et al.

1999; Ehrenfreund and Cami 2010, and references therein). These interstellar solids are nanoscale in dimensions, volatile in subsequent warmer solar nebular environments, and inaccessible from Earth; consequently, they will never grace the collections of mineralogical museums. Nevertheless, these ephemeral, sub-microscopic condensed phases are among the largest molecular repositories in the cosmos; thus, they play a significant role in the origins and distribution of key volatile molecular species in planetary systems.

At 10 K, most molecular species have a "sticking coefficient" close to unity, meaning that almost all gaseous species except for H_2 and He adhere to cold surfaces. Under such conditions, gradual molecule-by-molecule condensation of the most abundant molecular species, including H_2O, CO, CO_2, CH_3OH, and CH_4, is thought to occur on the surfaces of dust grains. By contrast, condensed molecular phases are not thought to be present in the warmer diffuse interstellar medium ($50 < T < 100$ K; $< 10^3$ atoms/cm^3; Ehrenfreund and Charnley 2000; Ferriere 2001).

The diversity and distribution of molecular components in interstellar solids reflect the initial gas phase composition, predominantly H and He with C, O, and N (Figure 1B), that has been significantly modified by several factors, including relative rates of condensation and sublimation at low temperatures, the nature of dust grain substrates, intermolecular chemical reactions, photochemical alteration by UV radiation and cosmic rays, and gradual annealing (Seki and Hasegawa 1983; Kouchi and Kuroda 1990; Jenniskens and Blake 1994, 1996; Kouchi et al. 1994; Ehrenfreund and Charnley 2000; Ehrenfreund and Fraser 2003; Gibb et al. 2004; Williams 2005; Hollenbach et al. 2009). For example, Allamandola et al. (1999) review the important role of hydrogen speciation, notably the ratio of H to H_2, in molecular clouds. In regions where monatomic hydrogen dominates, hydrogenation reactions lead to synthesis of H_2O, NH_3, and CH_4

297 as common condensed species. In H_2-rich environments, by contrast, O_2, N_2, CO, and CO_2 occur
298 more abundantly in condensates. Accordingly, Allamandola et al. (1999) conclude that condensed
299 molecular mantles on dust grains may fall into two principal populations distinguished by the
300 dominant local hydrogen species.
301
302 *Characterizing interstellar condensed phases*
303 The first interstellar molecular condensates to be identified were CO (Wilson et al. 1971) and
304 H_2O (Gillet and Forrest 1973), based on their strong absorption bands at 4.67 and 3.07 microns,
305 respectively. Subsequent discoveries of absorption and emission features point to hundreds of
306 molecular species, of which a few may condense and anneal in local concentrations that warrant
307 designation as an interstellar mineral. The most detailed picture of the nature and distribution of
308 these interstellar phases comes from orbiting telescopic measurements, for example as summarized
309 by Gibb et al. (2004), who review spectra in the 2.5- to 30-micron range measured by the Short-
310 Wavelength Spectrometer of the *Infrared Space Observatory (ISO)*. *ISO* spectra reveal diagnostic
311 absorption features from a range of molecular sources that point to more than a dozen forms of
312 interstellar condensed molecules, of which at least five—H_2O, CO, CO_2, CH_3OH, and CH_4—
313 occur as discrete nanoscale phases. Table 2 lists several phases that condense in the extremely
314 cold, low-pressure environments of circumstellar and interstellar molecular clouds.
315 By far the most abundant interstellar molecular condensate is water in its low-density
316 amorphous form (Jenniskens and Blake 1994, 1996), comprising more than 60% of observable
317 molecular solids (Gibb et al. 2004). Water molecules are identified based on a characteristic O-H
318 stretching infrared absorption feature at 3.05 microns, coupled with bending and combination

absorption modes at 1.65, 4.5, and 6.0 microns (Whittet 2003; Hagen et al. 1981; Grundy and Schmitt 1998; Newman et al. 2008). Importantly, interstellar water ice and other condensates were the primary source of water and other volatiles on Earth (Alexander et al. 2018).

Four other common condensate species are CO (with a diagnostic C-O stretch mode at 4.67 microns), CO_2 (C-O stretch and O-C-O bending modes at 4.27 and 15.2 microns, respectively), CH_3OH (four bands at 3.54, 3.95, 8.90, and 9.75 microns), and CH_4 (7.676 microns), all of which have been confirmed to exist in the solid state. These molecular solids, which formed primarily by heterogeneous condensation and/or photoreactions on dust grains (Seki and Hasegawa 1983; Lacy et al. 1991; Kouchi and Yamamoto 1995; Gibb et al. 2004; Williams 2005), occur both as relatively pure phases and as condensed molecular mixtures.

In addition to the five most abundant solid-forming molecules noted above, strong evidence exists for other molecular species as generally minor, i.e., < 1 molecular percent (mol %) components of condensed phases, including N_2, O_2, OCS, H_2CO, HCOOH, *X*CN (where *X* may be oxygen; Whittet et al. 2001), and NH_3, while the molecular ions OCN^-, NH_4^+, and CN^- may also play a role in some condensed solids (Gibb et al. 2004).

Molecular clouds also reveal complex assemblages of organic species, the majority of which accumulate in carbonaceous condensates analogous to soot and coal (Henning and Salama 1998; Ehrenfreund and Charnley 2000; Pendleton and Allamandola 2002; Snow and McCall 2006; Alexander et al. 2007; Tielens 2008). Ehrenfreund and Cami (2010) review carbon chemistry in the interstellar medium, in which almost 800 molecular species have been identified in molecular clouds (www.astrochemistry.net; accessed 12 August 2019), including significant quantities of nitriles, ketones, and esters (Allamandola et al. 1999; Yan et al. 2005; Ehrenfreund and Cami

2010). Of special interest are polycyclic aromatic hydrocarbons, or "PAHs," as well as a variety of fullerenes, both of which form in the high-temperature circumstellar envelopes of carbon-rich stars and have complex carbon chemistry analogous to soot formation in automobile exhaust or wood-fire smoke (Henning et al. 2004; Snow and Witt 1995; Ehrendfreund and Cami 2010; Salama et al. 2011).

SYSTEMATIC EVOLUTIONARY MINERALOGY: PART IIA—INTERSTELLAR MINERALOGY

From the mineralogical perspective, the nature of interstellar molecular solids presents challenges when attempting to define discrete phases. At temperatures close to 10 K, almost all molecular species will condense onto the cold surfaces of mineral grains immediately on contact (Greenberg 1991; Gibb et al. 2004). The initial molecular distribution, therefore, is highly disordered—a randomly condensed, amorphous molecular mixture rather than discrete, relatively pure phases. Only the most abundant molecular species are likely to gradually anneal into homogeneous nanoscale crystalline or amorphous volumes that might be justifiably characterized as mineral kinds. Such phases may be as small as 1 to 2 nanometers in diameter; experimental and theoretical experiments by Moberg et al. (2019) suggest that in the case of H_2O fewer than 100 molecules are required to form a discrete phase with the structural properties of ice.

Water ice, which accounts for 60 to 70% of molecular condensates in most observed molecular clouds (Whittet 2003; Gibb et al. 2004), provides the least ambiguous case. Specific sharp infrared emission features at 44 and 60 microns, as well as peak shapes and intensities of peaks at 1.65, ~3.1, and 4.53 microns, point unambiguously to both crystalline and amorphous condensed regions (Hagen et al. 1981; Moore and Hudson 1992; Grundy and Schmidt 1998; Newman et al. 2008). As noted above, at the relatively cold temperatures of dense molecular clouds, cubic ice is the expected crystalline form.

Carbon monoxide, often the second most abundant condensing molecular species at concentrations up to 20 % (e.g., Gibb et al. 2004), presents an intriguing case with two spectral types; the IR absorption spectra of solid CO differs significantly depending on its molecular environment (Sandford et al. 1988; Elsila et al. 1997; Williams 2005). For matrices composed

predominantly of polar species (predominantly H_2O), the ~4.67-micron peak has a broad component with a maximum at 4.682 microns. By contrast, a matrix of non-polar molecules (often CO, itself) is characterized by a narrower peak at 4.673 microns (Chiar et al. 1995, 1998). In the latter instance, a strong case can be made for discrete condensed amorphous CO as an interstellar phase.

Interstellar methanol is characterized by differences in absorption peak profiles that are also revealing, pointing to condensed regions of both relatively pure CH_3OH and methanol-water mixtures (Pontopiddan et al. 2003). However, for other less abundant interstellar molecules, of which hundreds have been catalogued (www.astrochemistry.net), the existence of discrete icy phases is inferred primarily on their relative abundances, rather than on specific spectroscopic characteristics.

Here we describe a few of the most abundant non-ionized interstellar molecules, including all C_1 molecules (those with only one carbon atom per molecule) and other species that are estimated to occur in some cool, dense molecular clouds at concentrations greater than 2 mol %. We catalog and name 8 condensed phases as likely interstellar minerals, while an additional 10 molecular species are labeled "unconfirmed as a discrete condensed phase" (Table 2).

Note that, with the exception of water ice, these 18 phases have not been identified in their crystalline forms naturally on Earth and are thus not in the current list of official IMA mineral species. Therefore, our nomenclature follows the conventions established by Hazen and Morrison (2020). The name of the molecule is preceded by the modifier interstellar (i.e., "*interstellar methane*") to distinguish it from similar species that may condense in circumstellar, nebular (see below), or planetary (Maynard-Casely et al. 2018; Hazen 2018) environments.

392

393 **NATIVE ELEMENTS**

394 *Interstellar hydrogen (H_2):* [unconfirmed as a discrete condensed phase] According to some
395 models of molecular condensation, a fraction of di-molecular hydrogen could condense at the
396 lowest temperatures (~10 K) in a cool dense molecular cloud. However, most hydrogen remains
397 in its gaseous form as H_2 or H atoms at T > 10 K.

398

399 *Interstellar nitrogen (N_2):* [unconfirmed as a discrete condensed phase] Molecular nitrogen
400 (N_2) is a relatively minor component of cool dense molecular clouds, but it should condense
401 heterogeneously with other molecules in environments with T < 20 K (Herbst and Klemperer 1973;
402 Womack et al. 1992; Knauth et al. 2004; Maret et al. 2006; Daranlot et al. 2012). In addition,
403 hydrogenation reactions of nitrogen play an important role in the formation of ammonia.

404

405 *Interstellar oxygen (O_2):* [unconfirmed as a discrete condensed phase] Oxygen is the third most
406 abundant element in most molecular clouds; nevertheless, molecular oxygen (O_2) is a minor
407 component of cool dense molecular clouds (Hollenbach et al. 2009; Wang et al. 2015). A
408 significant fraction of O_2 reacts with hydrogen to form water molecules.

409

410 **OXIDES**

411 *Interstellar cubic ice (H_2O):* Water (H_2O), both as an amorphous phase and crystalline ice, is
412 the most abundant interstellar condensed molecular species, comprising 60% or more of

413 condensates in cool dense molecular clouds (Whittet 2003; Gibb et al. 2004). The most familiar
414 crystalline structure of H_2O is the hexagonal form of snowflakes. However, below ~170 K the
415 stable crystalline form of ice condensed onto a cold substrate is "cubic ice" (Hobbs 1974). Because
416 water ice condenses at temperatures significantly higher than other common molecular solids (as
417 high as ~100 K), crystalline H_2O is also the only ice to commonly form mantles on silicate grains
418 as they are ejected from oxygen-rich stars (Kouchi and Yamamoto 1995), based on distinctive IR
419 absorption features (e.g., Whittet 2003).

420 "*Interstellar cubic ice (H_2O)*" in this evolutionary classification exclusively designates
421 crystalline H_2O, which is thought to account for significantly less than half of water condensates
422 in the interstellar medium. We retain the official IMA name, ice, for crystalline water, even though
423 the term "ice" is also commonly applied generically to other cold condensed molecular solids, both
424 crystalline and amorphous, in the cosmochemistry literature.
425
426 *Interstellar amorphous H_2O:* Amorphous H_2O is the most abundant interstellar mineral and is
427 a ubiquitous constituent of the icy mantles on presolar dust grains that form in cool dense molecular
428 clouds (Gibb et al. 2004). Amorphous H_2O may form by direct condensation at temperatures in a
429 molecular cloud, as well as by prolonged exposure of crystalline H_2O to ultraviolet or charged-
430 particle radiation (Kouchi and Kuroda 1990; Kouchi and Yamamoto 1995).
431 A complication is the possibility of multiple forms of amorphous H_2O (e.g., Kouchi 1987, 1990;
432 Sack and Baragiola 1993; Palumbo 2005), which depend on several factors, including the

433 temperature of condensation, the molecular/atomic flux density, exposure to ionizing radiation,
434 and the nature of the substrate.

435 The formation of amorphous versus crystalline H_2O is a balance between the temperature and
436 molecular/atomic flux. Crystalline water ices only form at relatively high temperature and low
437 molecular fluxes (e.g., < 10^8 molecules per cm^2 per second at 50 K; Kouchi et al. 1994). At the
438 lower temperatures and higher fluxes of cool, dense molecular clouds there is insufficient time for
439 molecular order to occur, so amorphous phases prevail. Gradual annealing to crystalline H_2O can
440 occur, but the process is highly temperature dependent and is balanced by gradual sublimation at
441 higher temperatures.

442

443 *Interstellar amorphous carbon monoxide (CO):* Interstellar carbon monoxide is a common
444 constituent of molecular clouds, identified by its prominent infrared absorption feature at 4.67
445 microns (Chiar et al. 1996; Elsila et al. 1997). Kouchi (1990) determined that CO condenses in
446 part as a separate amorphous phase and in part as an impurity up to a few percent in amorphous
447 H_2O (see also Collings et al. 2003). CO molecules can comprise a significant fraction (up to 25
448 mol %; Allamandola et al. 1999) of H_2O-dominated phases, with greater abundances in solids rich
449 in other nonpolar molecules, such as CO_2, O_2, and N_2 (Elsila et al 1997; Ehrenfreund et al. 1997).

450

451 *Interstellar carbon monoxide (CO):* Crystalline interstellar carbon monoxide has been shown
452 to form when amorphous CO warms to 23 K and anneals (Kouchi 1990).

453

454 *Interstellar amorphous carbon dioxide (CO_2):* Interstellar carbon dioxide has been identified
455 by its characteristic infrared absorption features at 4.27 and 15.2 microns as a ubiquitous
456 condensed species (d'Hendecourt and de Muizon 1989; Chiar et al. 1998; Gerakines et al. 1999;
457 Nummelin et al. 2001; Boogert and Ehrenfreund 2004; Gibb et al. 2004; Pontoppidan et al. 2008).
458 While carbon dioxide is relatively rare in the interstellar gas phase, it can represent as much as 20
459 mol % of a condensed molecular mixture in both single-phase form and in heterogeneous
460 molecular mixtures. This situation has led several authors to suggest that CO_2 forms principally
461 by oxidation of condensed CO in the solid state (Allamandola et al. 1999; Roser et al. 2001;
462 Williams 2005).
463
464 *Interstellar sulfur dioxide (SO_2):* [unconfirmed as a discrete condensed phase] Condensed
465 interstellar sulfur dioxide is revealed by absorption at ~7.6 microns (Boogert et al. 1997; Zasowski
466 et al. 2009).
467
468 **ORGANIC MOLECULAR SOLIDS**
469
470 *Interstellar amorphous methanol (CH_3OH):* The presence of interstellar methanol at
471 abundances as high as 25 mol % is revealed by a distinctive suite of infrared absorption bands
472 (Grim et al. 1991; Gibb et al. 2004). Pontopiddan et al. (2003) analyzed absorption peak profiles
473 that indicate regions of relatively pure CH_3OH, as well as methanol-water mixtures. The measured
474 methanol abundance in the interstellar medium is often significantly greater than that predicted
475 from models of gas-phase chemistry, possibly as a result of solid-state hydrogenation reactions in

476 mixed-phase water-CO condensed phases to make methanol (Williams 2005; Qasim et al. 2018)—
477 further evidence that the molecular compositions of interstellar solids are not exclusively a
478 consequence of condensation (Öberg et al. 2008).
479
480 *Interstellar amorphous methane (CH_4):* Interstellar methane at concentrations from ~1 to 4 mol
481 % is revealed by a diagnostic 7.676-micron absorption band (Lacy et al. 1991; Boogert et al. 1996;
482 d'Hendecourt et al. 1996; Gibb et al 2004). Spectroscopic evidence for a much lower concentration
483 of gas-phase methane suggests that CH_4 forms in the solid state through hydrogenation of atomic
484 carbon (Boogert et al. 1998). An unknown fraction of condensed interstellar methane may occur
485 in the form of clathrate hydrates (Ghose et al. 2019).
486
487 *Interstellar cyanide (XCN):* [unconfirmed as a discrete condensed phase] The presence of a
488 weak absorption feature at 4.62 microns in the spectra of some dense molecular clouds points to
489 the C-N stretch feature of cyanide, which may constitute up to a few mol % of some objects.
490 Generally given as "*X*CN," Whittet et al. (2001) suggest that *X* may be oxygen, though HCN is
491 likely present, as well (Snyder and Buhl 1971; Clark et al. 1974).
492
493 *Interstellar formaldehyde (H_2CO):* [unconfirmed as a discrete condensed phase] Diagnostic
494 microwave emissions at 4830 MHz (Snyder et al. 1969), as well as infrared features at 5.81 and
495 5.83 microns, point to molecular formaldehyde at abundances up to ~3 mol % (Grim et al. 1991;
496 Gibb et al. 2004).
497

498 *Interstellar formic acid (HCOOH):* [unconfirmed as a discrete condensed phase] The presence
499 of formic acid at concentrations less than 2 mol % is revealed by radio emissions (Zuckerman et
500 al. 1971; Winnewisser and Churchwell 1975), as well as by infrared absorption bands at 5.85 and
501 7.243 microns (Gibb et al. 2004). Peak profiles of IR bands reveal formic acid molecules in both
502 gas and condensed states (Schutte et al. 1999; Bisschop et al. 2007).
503
504 *Interstellar acetaldehyde (CH_3HCO):* [unconfirmed as a discrete condensed phase] An
505 absorption band at 7.414 microns suggests the presence of condensed acetaldehyde in significant
506 concentrations, perhaps as much as ~10 mol %, in many interstellar environments (Schutte et al.
507 1999; Gibb et al. 2004).
508
509 *Interstellar carbonyl sulfide (OCS):* [unconfirmed as a discrete condensed phase] Characterized
510 by a distinctive absorption feature at 4.91 microns, OCS is present at typically < 1 mol % (Palumbo
511 et al. 1995, 1997; Gibb et al. 2004).
512
513 *Interstellar ammonia (NH_3):* [unconfirmed as a discrete condensed phase] The presence of
514 interstellar ammonia is inferred from a minor absorption band at ~9 microns, close to the 9.3-
515 micron absorption of pure NH_3 (Smith et al. 1989; Chiar et al. 2000; Dartois et al. 2002). Note that
516 the position of the band shifts when the molecule is present in the polar environment of condensed
517 H_2O. This interpretation is reinforced by a band at 3.5 microns, which matches a broad feature
518 associated with ammonium hydrate (Gibb et al. 2004).
519

520 *Interstellar kerogen:* Laboratory studies of the effects of UV radiation on simple molecular
521 systems under high vacuum and low temperatures demonstrate the formation of complex insoluble
522 mixtures of non-volatile organic matter, not unlike the "kerogen" component that can be extracted
523 through acid dissolution of carbonaceous meteorites (Cronin and Pizzarello 1990; Greenberg et al.
524 1995; Widowiak et al. 1995; Ehrenfreund and Cami 2010). An absorption feature at ~3.47 microns
525 corresponds to the C-H stretch in hydrocarbons (Chiar et al. 1996), while several bands (3.3, 6.2,
526 8.6, and 11.3 microns) are consistent with polycyclic aromatic hydrocarbons – important carbon-
527 bearing components of carbonaceous meteorites (Grishko and Duley 2000). Kerridge (1983)
528 discovered isotopic heterogeneities in meteoritic kerogen consistent with low-temperature
529 interstellar origins. Ehrenfreund et al. (1991) demonstrated that distinctive features of the 3.4-
530 micron absorption band in the interstellar medium and in carbonaceous chondrites match closely.
531 Thus, it is likely that some, if not most, of the insoluble matter in the least altered (i.e., "3.0")
532 carbon-rich meteorites represents relatively unaltered interstellar material.
533 Note that in spite of their comprehensive list of meteorite phases, Rubin and Ma (2017) do not
534 include condensed organic material, i.e., "kerogen."

PRIMARY CONDENSATION MINERALOGY OF THE SOLAR NEBULA

Diverse and abundant chondrite meteorites provide an unparalleled view of the earliest stages of our solar system's formation (Brearley and Jones 1998; Krot et al. 2014; MacPherson 2014; Scott and Krot 2014; Connolly and Jones 2016). Data collected from these fascinating specimens, coupled with experimental measurements and theoretical modeling of high-temperature mineral condensation processes, provide an increasingly vivid picture of the first few million years of nebular evolution (e.g., Hashimoto 1983, 1992; Mysen et al. 1985; Mysen and Kushiro 1988; Davis et al. 1990; Nagahara et al. 1993; Wood and Hashimoto 1993; Ebel and Grossman 2000; Richter et al. 2002, 2007; Lodders 2003; Ebel 2006; Ebel and Alexander 2011; Davis and Richter 2014; Saxena and Hrubiak 2014; Han et al. 2015; Wood et al. 2019). The most primitive nebular phases survive in distinctive components of chondrite meteorites, including calcium-aluminum-rich inclusions (MacPherson 2014), amoeboid olivine aggregates (Grossman and Steele 1976; Krot et al. 2004), and ultra-refractory inclusions (El Goresy et al. 2002; Ma et al. 2014a). These ancient quasi-spheroidal to irregularly-shaped objects, ranging in size from less than 1 millimeter to more than a centimeter in diameter, have received extensive scrutiny and have been the subject of several comprehensive reviews (Brearley and Jones 1998; Ebel 2006; MacPherson 2014; Scott and Krot 2014; Rubin and Ma 2017, 2020). What follows, therefore, summarizes and codifies information that has been collated and reviewed by previous researchers.

Mineral evolution in the early solar nebula was driven by a succession of heating events, initially associated with the pre-main-sequence life of the Sun (Desch et al. 2012; Connolly and Jones 2016). Processes within this dynamic, evolving system are not fully understood, but include radiative heating, FU Orionis-type flares (Bertout 1989), bipolar outflows (Sahai et al. 2003), shock waves induced by infalling of gas (Iida et al. 2001), bow shocks from planetary embryos

(Desch and Connolly 2002; Boss and Durisen 2005) and associated magnetic effects (Mann et al. 2016; Mai et al. 2018), current sheets (Joung et al. 2004; Hubbard et al. 2012), and nebular lightning (Sorrell 1995; Desch and Cuzzi 2000), as well as cooling through thermal emission from the nebula. As larger objects grew through gravitational accretion, impact processes and magma formation increasingly played key roles. Therefore, primary mineralization in the early solar nebula occurred through condensation from a vapor enriched in O, Si, Mg, Al, Ca, and Fe (Figure 1C), followed by melt crystallization, and solid-state reactions. The formation of these primary nebular minerals has been the subject of extensive theoretical modeling (Urey 1955; Yoneda and Grossman 1995; Ebel and Grossman 2000; Ebel 2006), as well as experimental research under high-temperature, low-pressure conditions (Hashimoto 1983; Nagahara et al. 1993; Davis and Richter 2014).

The following section reviews the three most primitive types of chondrite objects, all of which contain primary condensation minerals of the early solar nebula: calcium-aluminum-rich inclusions, amoeboid olivine aggregates, and ultra-refractory inclusions (Figure 3).

In Part III of this series we will consider primary minerals in chondrules, which are abundant in many kinds of chondrite meteorites. Chondrules are igneous droplets thought to have formed between 1.5 and 3 million years after CAIs (e.g., Connolly and Jones 2016). Note that additional chondritic components, including opaque assemblages (also known as fremdlinge), dark inclusions, and fine-grained matrices, often contain similar suites of primary phases, but their histories are complicated by a succession of secondary processes associated with thermal, aqueous, and impact alteration (Brearley 1986; Blum et al. 1988; Scott et al. 1988; Johnson et al. 1990; Brearley and Jones 1998; MacPherson 2014; Rubin and Ma 2017). Thus, with the exception of

primary metal alloys, the mineralogy of these additional chondritic components will be considered in Parts IV and V.

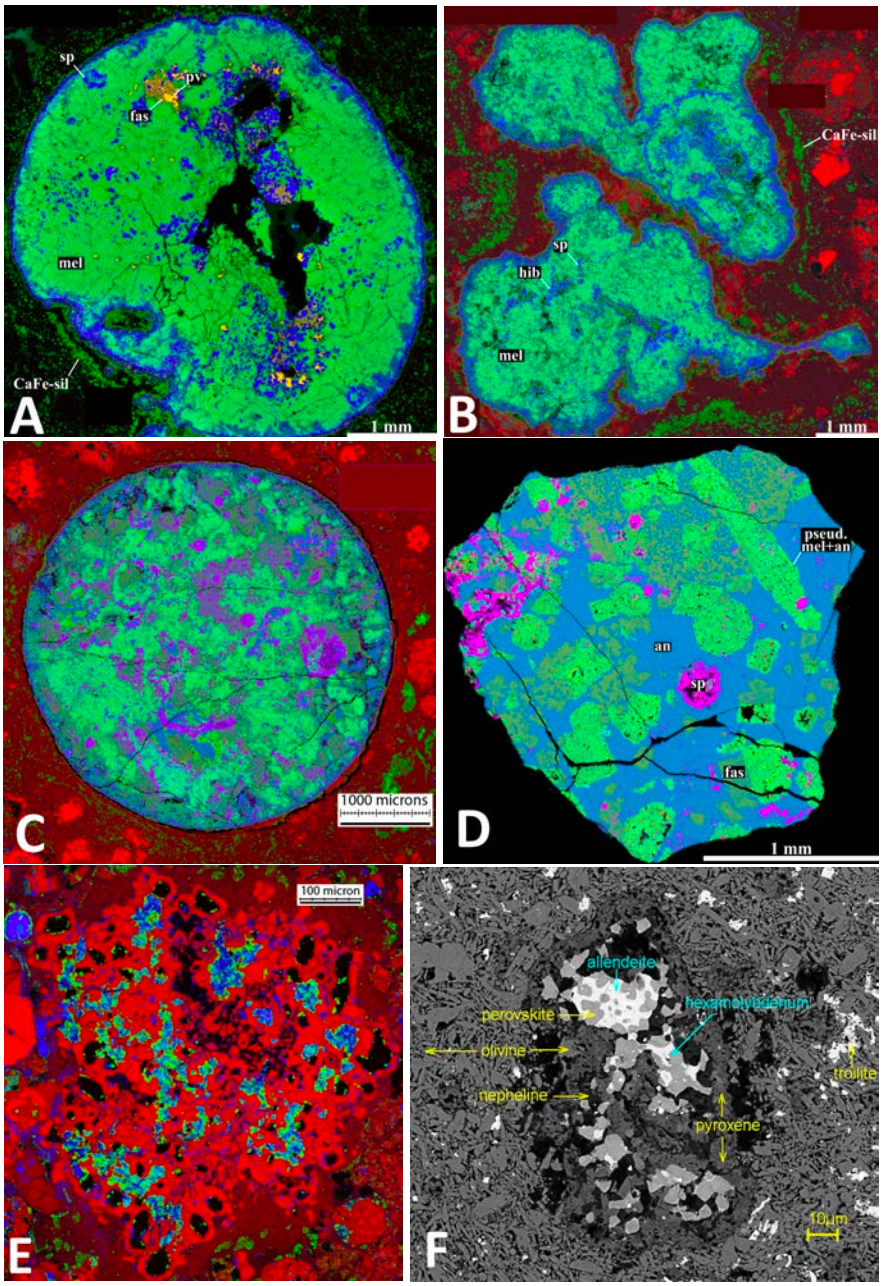


Figure 3. Components of primitive chondrite meteorites with primary mineral phases include: calcium-aluminum-rich inclusions (CAIs), amoeboid olivine aggregates (AOAs), and ultra-refractory inclusions (URIs). **A.** Compact type A CAI from the Adelaide meteorite with primarily mellilite and spinel and minor perovskite and fassaite (Ti = red; Ca = green; Al = blue); **B.** Fluffy type A CAI from the Allende meteorite with melilite, spinel, fassaite, and anorthite (Mg = red; Ca = green; Al = blue); **C.** Type B CAI from the Allende meteorite with melilite, fassaite, spinel, and anorthite (Mg = red; Ca = green; Al = blue); **D.** Type C CAI with anorthite, melilite, spinel, and

fassaite (Mg = red; Ca = green; Al = blue); **E.** Forsterite-rich AOA from the Kainsaz meteorite with dominant forsterite plus Fe-Ni metal alloys, fassaite, spinel, and anorthite (Mg = red; Ca = green; Al = blue); **F.** URI in matrix from the Allende carbonaceous chondrite. [Image credits: A, B, D, E courtesy of Alexander Krot, University of Hawaii; C courtesy of Denton Ebel, American Museum of Natural History; F courtesy of Chi Ma, Caltech]

Calcium-aluminum-rich inclusions (CAIs): CAIs are the most ancient objects formed in the solar nebula, with radiometric ages 4567.3 +/- 0.16 Ma (Amelin et al. 2002, 2010; Connelly et al. 2008, 2012; Krot 2019). These discrete components of primitive meteorites range in size from approximately 100 microns to more than a centimeter in diameter, and they possess distinctive suites of minerals reminiscent of ceramics—mineralogy that points to the earliest phase of solar nebular cooling and condensation (Brearley and Jones 1998; MacPherson 2014). Collectively, the refractory phases in CAIs represent the first ~5 percent of the solar nebula's condensable matter, as reflected in primitive CI chondrite meteorites (Grossman 1972; Davis and Richter 2014).

Detailed observations of CAIs from diverse types of chondrite meteorites reveal significant variations in bulk composition, mineralogy, and morphology—differences that reflect nebular heterogeneities in space and time (Liu et al. 2009; Kööp et al. 2018; MacPherson 2014 and references therein), as well as subsequent processing through re-melting, distillation, and a variety of secondary processes (the subject of Parts IV and V of this series). Owing to their diversity, CAIs have been divided into several groups according to size, shape, and mineralogy (Grossman 1975; Wark 1987; Beckett and Stolper 1994; MacPherson 2014):

- *Compact Type A CAIs* are spheroidal objects, evidently a consequence of partial melting. Melilite is often the dominant mineral with spinel, hibonite, perovskite, and other refractory phases (Figure 3A).
- *Fluffy Type A CAIs* have similar mineralogy to compact type A CAIs but their shapes are irregular, suggesting minimal melting (Figure 3B).

- *Type B CAIs* are spheroidal objects characterized by their relatively large size (up to several centimeters) and mineralogy that commonly includes melilite, spinel, anorthite, forsterite, and the Al-Ti-rich calcic clinopyroxene called fassaite (Figure 3C). Type B CAIs are further divided into subcategories based on mineralogical details.
- *Type C CAIs* consist primarily of spinel, fassaite, and anorthite with textures characteristic of igneous processes (Figure 3D).
- In addition to A-, B-, and C-type CAIs, researchers recognize other fine-grained (i.e., < 200 microns) objects with CAI affinities. *Spinel-hibonite-rich spherules* (sometimes abbreviated SHIBs) are CAIs that contain a relatively unaltered assemblage of spinel, hibonite, and perovskite, with associated melilite, fassaite, and anorthite (Ireland 1988; Kööp et al. 2016a).
- *Platy hibonite crystals*, known as PLACs, are CAIs that consist of lath-like crystals of hibonite, sometimes surrounded by a silicate glass or fassaite, often with hibonite, grossite, and melilite, but typically lacking spinel (Ireland 1988; Kööp et al. 2016b).
- A mineralogically curious type of CAI, dubbed *"FUN" CAIs* by Wasserburg et al. (1977), displays "Fractionation and Unidentified Nuclear effects" – specifically, large mass-dependent fractionations of Mg, Si, and O, as well as other isotopic anomalies. In terms of primary mineralogy FUN CAIs are, for the most part, similar to type B CAIs. However, the discovery of dmisteinbergite ($CaAl_2Si_2O_8$), a rare hexagonal high-temperature polymorph of anorthite, points to an unusual formation environment close to the protosun with T > 1500 K and P < 10^{-6} (i.e., significantly lower pressure than the ~10^{-4} atm modeled for most other CAIs; Ma et al. 2013a).

In our review we focus on the primary mineralogy of CAIs and do not make distinctions among these varieties of calcium-aluminum-rich inclusions. Note, however, that we do distinguish CAIs from ultra-refractory inclusions, as the latter began to form later and are chemically and mineralogically distinct from CAIs.

Isotopic evidence, including surprisingly uniform anomalous Mg and O isotopes across all types of CAIs (McKeegan et al. 1998; MacPherson 2014 and references therein) and age determinations by various methods (Connelly et al. 2012; see also, Amelin et al. 2002, 2010; Connelly et al. 2008; Krot 2019), reveal that CAIs formed at ~4567.3 +/- 0.16 Ga over a span of ~200,000 years. However, detailed chemical and isotopic investigations point to at least two stages of CAI formation as a consequence of rapid protosun evolution during the earliest stages of the solar system.

The earliest nebular condensates are represented by a relatively small population of platy hibonite-bearing "PLAC" CAIs, which appear to have formed prior to an influx of ^{26}Al to the solar nebula, as measured by diagnostic ^{26}Mg (Liu et al. 2009). Hibonite in these CAIs contains helium and neon excesses that could only have formed under an intense flux of energetic particles, albeit at a heliocentric distance significantly beyond the inner disk edge, where temperatures were cooler (Kööp et al. 2018).

A second generation of CAIs of the spinel-hibonite-bearing "SHIB" type display at least three important differences from PLAC CAIs: (1) they contain excess ^{10}B formed by the decay of short-lived ^{10}Be, which points to condensation near the inner edge of the nebular disk in a single zone close to the Sun (McKeegan et al. 2000; MacPherson et al. 2003; Krot 2019); (2) they incorporated significant short-lived ^{26}Al (preserved as diagnostic ^{26}Mg), pointing to formation after PLAC

CAIs (Kööp et al. 2016a, 2016b); and (3) they lack the anomalous helium and neon contents of PLAC hibonite (Kööp et al. 2018). According to current models, these CAIs were later dispersed to beyond Jupiter's orbit by strong solar winds (Shu et al. 1996). Further studies of CAI isotopic and mineralogical variations are likely to refine our understanding of the earliest stages of solar system evolution.

CAIs rarely represent unaltered primordial nebular condensates. They typically have been subjected to multiple stages of partial alteration from reheating, including annealing, melting, and distillation; impact melting, volatilization, and metamorphism; and secondary alteration, including oxidation, sulfidization, and hydration (MacPherson and Davis 1993; Beckett et al. 2000; Rubin and Ma 2017). Here we focus exclusively on the so-called "primary" mineralization in the sense of MacPherson (2014) – i.e., mineralization that results from "direct condensation, melt solidification, or solid-state recrystallization." Note that in this context solid-state recrystallization includes only the first phase of reaction and replacement of earlier high-temperature phases by exsolution, element ordering, or reversible phase transition, during equilibrium cooling of CAIs below their initial condensation temperature and prior to their incorporation into larger bodies (e.g., Yoneda and Grossman 1995; Ebel 2006). Chondrite mineralization through subsequent secondary processes, including oxidation, sulfidation, aqueous and thermal alteration (metamorphism), shock processes, and planetesimal differentiation, will be considered in Parts IV and V of this series.

Only seven minerals, all of which occur in calculations of nebular condensation sequences between ~1800 and 1100 K (e.g., Grossman 1972; Brearley and Jones 1988; Lodders 2006; Davis and Richter 2014; MacPherson 2014, Table 1; Wood et al. 2019), are common primary minerals in CAIs. In order of appearance, they are hibonite, perovskite, gehlenite/åkermanite (both members

684 of the melilite group), spinel, Al-Ti-rich calcic clinopyroxene ("fassaite"), forsterite, and anorthite.
685 Note that corundum is calculated to be the highest-temperature oxide condensate (Lattimer et al.
686 1978; Anders and Grevesse 1989; Lodders 2003; Wood et al. 2019), yet it is usually absent owing
687 to subsequent reactions with nebular vapor to hibonite, spinel, or melilite. Additional phases,
688 though less common, include the oxides grossite, krotite, panguite, and tistarite; the calcium-
689 titanium silicate rhönite; and alloys of Fe-Ni and platinum group elements (see Table 3).
690 An unresolved question related to CAIs is the paragenesis of ubiquitous thin outer layers, first
691 described by Wark and Lovering (1977) and now known as "Wark-Lovering rims" (MacPherson
692 et al. 1981; Ruzicka 1997; Wark and Boynton 2001). These layers are typically a few tens of
693 microns thick and consist of both refractory condensates, often spinel, melilite, and fassaite,
694 occasionally with hibonite, perovskite, anorthite, and forsterite, as well as such secondary minerals
695 as nepheline, Fe-rich spinel, hedenbergite, and andradite (MacPherson 2014, and references
696 therein). No new obviously primary minerals occur in these rims; therefore, for the purposes of
697 this review, even though they may represent a distinct timing and mode of mineral formation, we
698 do not consider Wark-Lovering rims as representing distinct natural kinds. Note, however, an
699 important feature of this evolutionary system of mineralogy is that specialists have the option of
700 splitting mineral natural kinds into finer and finer sub-categories, based on their distinctive modes
701 of origin, as reflected in diagnostic suites of physical and chemical attributes.
702
703 *Amoeboid olivine aggregates (AOAs):* Amoeboid olivine aggregates (AOAs) represent a
704 second common refractory constituent of chondrite meteorites. Originally described from the
705 Allende CV chondrite by Grossman and Steele (1976) and subsequently identified from a wide
706 variety of carbonaceous meteorites (Grossman et al. 1979; Kornacki and Wood 1984; McPherson

et al. 1988; Aléon et al. 2002; Chizmadia et al. 2002; Krot et al. 2004; Rubin 2013), AOAs are thought to have formed after CAIs, though before chondrules that were forming at ~ 4565 +/- 0.5 Ma (Connolly and Jones 2016), and at generally lower condensation temperatures and at pressures consistent with 10^{-3} atm (Weisberg et al. 2004). They typically occur as irregularly-shaped assemblages up to 0.5 millimeters in diameter, constituting a few percent of some carbonaceous chondrites (Scott and Krot 2014; Figure 3E). AOAs consist of sintered accumulations of nebular condensates, primarily forsterite (typically fine-grained, < 20 microns), with Fe-Ni metal alloys and a refractory assemblage, commonly including fassaite, spinel, and anorthite, and occasionally with perovskite and/or melilite. Models of AOA evolution suggest that other refractory phases, including corundum, grossite, and hibonite, may have initially formed but were transformed to melilite, clinopyroxene, spinel, and/or anorthite through continuous solid-state reactions of the fine-grained constituents on cooling (Weisberg et al. 2004).

Many AOAs also hold a rich variety of secondary minerals, including phyllosilicates, feldspathoids, sulfides, and other phases, which will be considered in Part V.

Ultra-Refractory Inclusions (URIs): Ultra-refractory inclusions (URIs) are a scarce population of mineralogically important pre-chondrule objects (Figure 3F), usually grouped with CAIs, that are extremely enriched in Sc, Zr, Y, and other refractory elements by factors of as much as 10^3 (El Goresy et al. 2002). A consequence of these enrichments is the appearance of more than a dozen rare oxide and silicate minerals (Rubin and Ma 2017), including allendeite ($Sc_4Zr_3O_{12}$), thortveitite ($Sc_2Si_2O_7$), and lakargite ($CaZrO_3$), as well as Sc- and Ti^{3+}-bearing clinopyroxenes and garnets (Table 3). The unusual chemistry of URIs points to an early stage of nebular evolution

729 in regions that had already generated significant quantities of primary CAI minerals with Mg, Ca,
730 Al, and Ti. URIs are often found as inclusions in later chondrite objects, including AOAs,
731 chondrules, and fine-grained matrix (El Goresy et al. 2002; Ma et al. 2009a, 2014a).
732
733 *A note regarding opaque assemblages*
734 Some researchers have suggested that Fe-Ni-metal-rich opaque assemblages (also referred to
735 as "fremdlinge"), which are found associated with some CAIs (e.g., Brearley and Jones 1998;
736 MacPherson 2014), represent some of the earliest primary condensates of the solar nebula (El
737 Goresy et al. 1978; Armstrong et al. 1985, 1987). Others posit a later origin of these objects with
738 a complex history of alteration, for example by oxidation, sulfidation, and exsolution through
739 exposure to heterogeneous nebular environments (Wark and Lovering 1982a; Blum et al. 1988,
740 1989; MacPherson 2014). We include the primary metals of opaque assemblages – Fe-Ni and
741 platinum group element alloys – in the list of CAI minerals. However, the lower-temperature
742 mineral suites of opaque assemblage minerals, including fayalite, wollastonite, feldspathoids,
743 tungstate-molybdates, and varied sulfides and phosphates, will be considered in Parts IV and V.

SYSTEMATIC EVOLUTIONARY MINERALOGY: PART IIB—PRIMARY NEBULAR MINERALOGY

In the following section we outline the mineralogy of primary nebular phases, including minerals in CAIs, AOAs, and URIs formed by condensation, melt crystallization, and their initial solid-state reactions (see also Table 3). We tabulate 59 natural kinds of primary minerals, corresponding to more than 40 IMA-approved species plus nebular silicate glass. We also include one low-temperature solar nebular condensate, cubic H_2O ice, which has been observed through telescopic observations in the cool (~ 100 K) outer regions of other stellar environments. Each mineral natural kind is given a binomial designation: the first name indicates the paragenetic context (e.g., *CAI*, *AOA*, or *URI*), whereas the second name for the most part conforms to approved IMA mineral species' names. However, in several instances we deviate from IMA conventions:

- *Refractory metal alloys:* CAIs, URIs, and associated opaque assemblages often contain micron-scale "nuggets" in which refractory metals such as Mo, Ir, Os, Ru, Rh, Re, Pt, W, and Ru occur in hexagonal ($P6_3/mmc$) alloys with widely varied elemental proportions (e.g., Weber and Bischoff 1997; Berg et al. 2009; MacPherson 2014; Chi Ma, personal communication). For example, Berg et al. (2009) report the compositions of 88 refractory metal nuggets from the Murchison meteorite, most with significant concentrations of Mo, Os, and Ru, with each of those elements dominant in some grains. IMA protocols require naming each alloy based on the most abundant element; therefore, ruthenium, osmium, molybdenum, and other metals have all been officially recognized as separate native element mineral species. However, because these alloys form continuous solid solutions in chondrite nuggets and they are all formed by the same paragenetic process of condensation from a high-temperature nebular gas, we lump them together into platinum group element

(PGE) alloys. Similarly, we lump the cubic (*Fm*3*m*) iron-molybdenum alloys hexaferrum (Fe,Os,Ir,Mo) and hexamolybdenum (Mo,Ru,Fe) into one natural kind: Fe-Mo alloys.

- *Anosovite:* The name "anosovite" for pseudobrookite-type Ti_3O_5 was discredited by Bowles (1988), as it was only known as an anthropogenic phase in Ti-rich slags. However, Zhang et al. (2015) have discovered this phase as a primary nebular condensate and we resurrect the name anosovite, pending IMA's decision on the naming of this mineral.
- *Melilite:* In the case of the melilite group, a complete solid solution exists between gehlenite ($Ca_2Al_2SiO_7$) and åkermanite ($Ca_2MgSi_2O_7$), with primary meteoritic compositions ranging from $Åk_{01}$ to $Åk_{100}$ (Brearley and Jones 1998). Individual grains, furthermore, may be zoned such that cores are technically gehlenite and rims technically åkermanite. These variations occur within a continuous solid solution and represent a single mode of formation. Therefore, we classify melilite samples as "*CAI melilite*," "*AOA melilite*," or "*URI melilite*."
- *Fassaite:* A fourth deviation from standard IMA nomenclature relates to primary CAI, AOA, and URI pyroxenes, typically a Ca-Mg-dominant, Fe-poor clinopyroxene with significant Al and Ti (both Ti^{3+} and Ti^{4+}) in solid solution. The IMA-approved name for most of these occurrences is diopside, because the closest compositional end-member is $CaMgSi_2O_6$. However, the great majority of near-end-member diopside occurrences in CAIs are of secondary origin (Brearley and Jones 1998), and thus should be distinguished from the primary Al-Ti-bearing nebular calcic clinopyroxene phase, $[Ca(Mg,Al,Ti^{3+},Ti^{4+})(Al,Si)SiO_6]$. This complex solid solution has long been called

"fassaite" in the meteoritics literature, based on its similarity to iron-poor calcic clinopyroxenes from the Fassa Valley, Trento Province, Italy (Thompson 1818; Deer et al. 1963; Sack and Ghiorso 2017). The name fassaite was discredited during a reclassification of pyroxene nomenclature (Morimoto et al. 1988; see also Hazen 1989), but it continues to be used by many meteorite experts to describe the distinctive primary calcic clinopyroxenes in chondrite meteorites (e.g., Brearley and Jones 1998; MacPherson 2014; Rubin and Ma 2017; Sack and Ghiorso 2017). Primary clinopyroxene from meteorites possesses a number of diagnostic attributes, including extensive Al, Ti^{3+}, Ti^{4+}, and (less commonly) Sc and/or V solid solution – compositional complexities that lead to a suite of distinctive optical properties, including refractive indices, birefringence, extinction angle, optic axial angle, and dispersion (e.g., Deer et al. 1963). Therefore, we retain the name fassaite for Al-Ti-bearing clinopyroxene.

- *Rhönite:* Finally, we lump two rare primary CAI minerals of the sapphirine group, rhönite [$Ca_2(Mg,Al,Ti)_6(Si,Al)_6O_{20}$] and addibischoffite [$Ca_2(Al,Mg,V,Ti)_6(Al,Si)_6O_{20}$). These similar phases represent a continuous solid solution and occur via the same paragenetic mode.

NATIVE ELEMENTS

Iron-nickel alloys are common primary and secondary phases associated with CAIs, AOAs, URIs, and associated opaque aggregates (Brearley and Jones 1998; Rubin and Ma 2020). Chondrites also often incorporate micron-scale nuggets of highly refractory metal alloys, which contain Ir, Os, Ru, Mo, and other siderophile elements (Sylvester et al. 1993; Berg et al. 2009). In

addition, refractory mineral assemblages occasionally incorporate metal grains with significant amounts of both Fe and Mo (Ma et al. 2014a; Zhang et al. 2015; Rubin and Ma 2017).

Platinum Group Element (PGE) Alloys (Os,Ir,Ru,Rh,Pt,W,Mo,Re): CAIs often incorporate micron-scale "refractory metal nuggets" (hexagonal, $P6_3/mmc$), containing elements of the platinum group (Os, Ir, Ru, Rh, and Pt), as well as Mo, W, and Re (Palme et al. 1994; Brearley and Jones 1998; MacPherson 2014), which condensed at temperatures between 1800 and 1300 K at 10^{-4} atm (Berg et al. 2009; Harries et al. 2012; Scott and Krot 2014). Some individual sub-micron-scale grains are close to pure Pt, Ru, or Re, whereas others are multi-element alloys with Os, Ir, Ru, or Pt as the most abundant metal (e.g., El Goresy et al. 1978; Wark and Lovering, 1978; Bischoff and Palme, 1987; Brearley and Jones 1998; Berg et al. 2009). Because these elements form extensive solid solutions by the same nebular condensation mechanism, we lump them together as "*PGE alloys*."

CAI PGE alloys: Refractory metal nugget alloys of PGEs plus W, Mo, and Re are commonly found as micron-scale grains, often as inclusions in oxides or silicates, in CAIs (e.g., El Goresy et al. 1978, 1979, 1984; Wark and Lovering 1982b; Sylvester et al. 1993; Endress et al. 1994; Geiger and Bischoff 1995; Weber and Bischoff 1997; Berg et al. 2009; MacPherson 2014).

URI PGE alloys: Micron-scale grains of Os-dominant PGE alloys are found in ultra-refractory inclusions in association with typical CAI phases, including fassaite (often Sc- and Ti-rich), perovskite, and spinel (Ma and Rossman 2008; Ma 2011; Ma et al. 2014a).

832 **Iron-Nickel Alloys:** Iron-nickel alloys condensed from the solar nebula at temperatures
833 estimated between 1350 and 1450 K (Campbell et al. 2005). Consequently, metal alloys with iron
834 dominant, typically incorporating significant Ni and at times with minor amounts of other
835 siderophile elements, are common as a minor phase in CAIs, AOAs, and URIs from many types
836 of chondrites. They occur as the minerals iron (also known as "kamacite"), taenite, and awaruite.
837 Here we accept iron and taenite as primary nebular condensates. Note that awaruite, an
838 isometric (*Pm*3*m*) Ni-dominant alloy of iron and nickel, is also a common minor metallic phase in
839 opaque aggregates associated with CAIs (Taylor et al. 1981; Rubin and Kallemeyn 1989; Ikeda
840 1992; Smith et al. 1993; Casanova and Simon 1994; Moggi-Cecchi et al. 2007). However, it
841 appears that most occurrences of awaruite in CAIs are of secondary origin (Brearley and Jones
842 1998; Rubin and Ma 2020).
843
844 **Iron (alpha-Fe):** Native iron, also known as kamacite (a discredited though often used mineral
845 name in the context of meteorites), is the most stable low-Ni alloy of Fe and Ni. This cubic (*Im*3*m*)
846 phase has nickel contents that are typically less than 10 wt. % Ni, while Co is less than 1 wt. %
847 (Brearley and Jones 1998). Native iron is common as an opaque phase associated with CAIs,
848 AOAs, and URIs (Bevan and Axon 1980; Zinner et al. 1991; Simon and Grossman 1992; Caillet
849 et al. 1993; Zhang et al. 1995; Shibata 1996; Ma and Rossman 2008).
850
851 *CAI iron:* Native iron has been reported as a primary phase (Campbell et al. 2005; Scott and
852 Krot 2014; Rubin and Ma 2017) in type A (MacPherson and Grossman 1984; Ulyanov et al. 1982),
853 type B (Blander and Fuchs 1975; Sylvester et al. 1992; Caillet et al. 1993), and rarely in type C
854 CAIs (Blander et al. 1980).

855

856 *AOA iron:* Iron-nickel alloys, typically with 5 to 7 wt. % Ni, are ubiquitous components of
857 amoeboid olivine aggregates (Weisberg et al. 1993, 2004; Chizmadia et al. 2002; Krot et al. 2004).
858 They occur as blebs up to 10 microns in diameter, often as inclusions in refractory oxides and
859 silicates.

860

861 *URI iron:* Ma and Rossman (2008) recorded sub-micron grains of Fe-Ni alloy associated with
862 zirconolite, tazheranite, and PGE alloys in an ultra-refractory inclusion from the Allende chondrite.

863

864 **Taenite [gamma-(Fe,Ni)]:** Taenite is an isometric (*Pm*3*m*) alloy of gamma-iron, typically with
865 10 to more than 50 wt. % Ni (Affiatalab and Wasson 1980; Nagahara 1982). Taenite often occurs
866 in CAIs and opaque assemblages as a minor phase, typically in close association with kamacite
867 (Zinner et al. 1991; Simon and Grossman 1992; Sylvester et al. 1992; Caillet et al. 1993; Brearley
868 and Jones 1998), in some instances as exsolution lamellae in kamacite (Noguchi 1994; Ichikawa
869 and Ikeda 1995). Taenite and kamacite are also associated in fine-grained mixtures known as
870 "plessite" (Massalski et al. 1966; Scott and Rajan 1979).

871

872 *CAI taenite:* Taenite commonly occurs, both as isolated grains and as exsolution lamellae in
873 kamacite, in types A and B CAIs (MacPherson and Davis 1993; Ichikawa and Ikeda 1995;
874 Campbell et al. 2005; Rubin and Ma 2017).

875

876 **Iron-Molybdenum Alloys (Fe,Mo):** Iron and molybdenum, in combination with refractory
877 metals, notably Ir, Os, and Ru, occasionally form micron-scale nuggets of hexagonal ($P6_3/mmc$)

878 alloys, which must have condensed or solidified at temperatures intermediate between the ultra-
879 refractory platinum group alloys and the lower-temperature Fe-Ni alloys described above. At least
880 two IMA-approved species have been identified from CAIs in the Allende meteorite – hexaferrum
881 [(Fe,Os,Ir,Mo); Ma 2012; Zhang et al. 2015] and hexamolybdenum [(Mo,Ru,Fe); Ma et al. 2011b,
882 2014a]. Ma et al. (2014a) provide a review of this "continuum of meteoritic refractory alloys with
883 the $P6_3/mmc$ structure."
884
885 *CAI Fe-Mo alloys:* Ma et al. (2011a) and Ma (2012) described micron-scale nuggets of
886 hexagonal alloys with varying proportions of Fe, Mo, Ru, and other metals from CAIs in the
887 Allende and NWA 1934 meteorites.
888
889 *URI Fe-Mo alloys:* Ma et al. (2014a) found hexamolybdenum in association with allendeite in
890 an ultra-refractory inclusion, while Zhang et al. (2015) report hexaferrum.
891
892 **CARBIDES**
893 Khamrabaevite (TiC) is the only confirmed refractory carbide with characteristics of primary
894 nebular phases (Ma and Rossman 2009a). Iron carbide minerals, including cohenite [$(Fe,Ni)_3C$],
895 edscottite (Fe_5C_2), and haxonite [$(Fe,Ni)_{23}C_6$], are known from highly reduced chondrules in
896 carbonaceous chondrites and enstatite chondrites (e.g.,; MacPherson 2014; Rubin and Ma 2017;
897 Ma and Rubin 2019), but they have not been documented as primary phases in CAIs, AOAs, or
898 URIs. Consequently, they are considered in Part III of this series.
899

URI khamrabaevite: Titanium carbide (TiC) was identified by Ma and Rossman (2009c) as a 10-micron diameter grain associated with corundum and tistarite from the Allende carbonaceous chondrite.

NITRIDES

Nitrides are rare in meteorites. Nierite (Si_3N_4) is known as a stellar mineral (Hazen and Morrison 2020), but has not to our knowledge been reported in CAIs, AOAs, or URIs. Sinoite (Si_2N_2O) is known as both a primary chondrule mineral and as an impact product in enstatite chondrites (Lin et al. 2011; El Goresy et al. 2011; see Parts III and IV).

The only known candidate for a primary nebular condensate mineral is osbornite (TiN), which is known from chondrules of enstatite chondrites (e.g., El Goresy et al. 2011), as well as from CAIs in several carbonaceous chondrites (Weisberg et al. 1988; Grokhovsky 2006; Krot et al. 2006; MacPherson 2014). Note that osbornite is also known as a secondary mineral in enstatite chondrite impact melts (Rubin and Ma 2020).

CAI osbornite: Osbornite (TiN) was identified in CAIs from the Isheyevo (Grokhovsky 2006; Krot et al. 2006) and Allan Hills 85085 (Weisberg et al. 1988) carbonaceous chondrites. Osbornite was also found in Stardust samples that were presumed to be CAI fragments (Weisberg et al. 2006).

SILICIDES

921 The silicide perryite [$(Ni,Fe)_5(Si,P)_2$] is a refractory phase with characteristics of primary
922 condensates (MacPherson 2014; Rubin and Ma 2017). However, perryite is found principally in
923 the highly reduced mineral assemblages of enstatite chondrite chondrules and will be described in
924 Part III of this series.
925
926 **PHOSPHIDES**
927 At least two phosphides, monipite (MoNiP) and schreibersite [$(Fe,Ni)_3P$], have been described
928 as primary phases in chondrite meteorites. The exact mode of formation of these grains is
929 uncertain; plausible hypotheses include reaction of a P-bearing nebular gas-phase with an Fe-Ni
930 alloy (Schaefer and Fegley 2010), crystallization from a P-rich immiscible melt that exsolved from
931 Fe-Ni melt, or exsolution from a solidified Fe-Ni alloy (Ma et al. 2014b). In this tabulation we list
932 only monipite, which was discovered in a CAI. Schreibersite is known principally from highly
933 reduced assemblages in enstatite chondrite chondrules and will be considered in Part III.
934
935 *CAI monipite:* Monipite with composition [$(Mo_{0.84}Fe_{0.06}Co_{0.04}Rh_{0.03})(Ni_{0.89}Ru_{0.09})P$] was
936 described by Ma et al. (2014b) from a 1- x 2-micron crystal in a type B CAI from the Allende
937 meteorite. It occurs in association with primary phases melilite, fassaite, and spinel, as well as
938 probable alteration minerals, including awaruite, the rare oxides kamiokite ($Fe_2Mo_3O_8$),
939 tugarinovite (MoO_2), and an unnamed Nb-rich oxide [$(Nb,V,Fe)O_2$].
940
941 **SULFIDES**

A number of sulfide minerals, including niningerite (MgS), oldhamite (CaS), pentlandite [$(Fe,Ni)_9S_8$], and troilite (FeS), probably formed early in the history of the solar nebula (Rubin and Ma 2017, 2020). However, all of these phases, as well as numerous subsequent sulfides, either formed by igneous processes in chondrules (e.g., Rubin et al. 1999; MacPherson 2014; see Part III), by sulfidation of earlier phases through reaction with an S-rich vapor, or by solid-state reactions (Rubin and Ma 2017). Therefore, we do not list any sulfide minerals as primary nebular condensates in Part II.

Troilite represents a difficult case. In some instances, it appears to be a primary chondrule mineral (e.g., Rubin et al. 1999; El Goresy et al. 2011) and thus will be included in Part III of this series. However, many occurrences in meteorites, including in CAIs of enstatite chondrites, appear to represent secondary mineralization by sulfidation (Fagan et al. 2000; Guan et al. 2000; MacPherson 2014). Therefore, in spite of instances where troilite in CAIs is associated with primary oxide and silicate phases and a primary origin for the sulfide cannot be ruled out (Fagan et al. 2001), we do not list any sulfides as primary nebular condensates in CAIs, AOAs, or URIs.

OXIDES

Oxygen is the most abundant element in the solar nebula after hydrogen and helium, and it played a dominant role in the condensation of primary refractory phases in CAIs, AOAs, and URIs. The most common primary nebular oxides in CAIs, AOAs, and URIs contain Mg, Ca, Al, and/or Ti, with rare minor oxides of Mo, Sc, V, and Zr. We also include ice (cubic H_2O) as the only molecular crystal likely to have condensed in the cool (~ 100K) outer regions of the solar nebula.

Note that more than two dozen other refractory oxides are recorded as primary minerals from chondrules (MacPherson 2014; Rubin and Ma 2017). The distinction between primary condensates

in CAIs, AOAs, and URIs (the subject of this contribution) versus primary igneous phases in chondrules (which will be summarized in Part III) is important. The limited number of confirmed CAI, AOA, and URI oxide phases listed below are thought to have formed by condensation via cooling of high-temperature, low-pressure vapor or their subsequent solid-state reactions in the earliest solar nebula at ~4.567 Ga. Subsequent primary phases in chondrules were formed by partial melting of nebular materials at least 1.5 million years later – a time when planetesimal formation and nebular heterogeneities had become well established (e.g., Gilmour and Saxton 2001; Connolly and Jones 2016). Primary chondrule minerals thus crystallized at a later time by a different combination of processes; i.e., at generally lower temperatures, during relatively short heating events, and on nebular material that had already undergone a preliminary stage of chemical and isotopic fractionation.

Confusion can arise as to what constitutes a primary CAI, AOA, or URI phase. For example, the rare mineral tistarite (Ti_2O_3) was described from a single 1- x 2-micron grain in a chondrule from the Allende carbonaceous chondrite (Ma and Rossman 2009c). Because of its presumed primordial character, MacPherson (2014, Table 1) listed tistarite as one of only 15 CAI primary minerals. However, coexisting rutile (TiO_2) is not included in this tabulation of the earliest primary nebular minerals, because it appears to represent a later stage of nebular mineralization (Ma 2019). Numerous other meteoritic oxide minerals will be considered in Parts III, IV, and V of this series.

Ice (H_2O): Primary solar nebular mineralogy consists almost exclusively of high-temperature (T > 1100 K) refractory phases that form by condensation from a vapor phase or melt crystallization in close proximity to the central star. However, infrared telescopic observations

reveal the presence of crystalline H_2O, presumably condensed as thin mantles (< 0.05 microns thick) on oxide and silicate dust grains in the cold circumstellar regions of some oxygen-rich stars (Omont et al. 1990). We therefore include ice as a primary nebular condensate.

The distinction between crystalline and amorphous condensed H_2O in the context of circumstellar environments is important. Crystalline water, which is revealed by sharp IR emission features at 44 and 60 microns, can only condense directly at temperatures > 50 K under relatively low molecular fluxes (estimated at $\sim 10^4$ molecules per cm^2 per second in circumstellar environments; Kouchi and Yamamoto 1995). By contrast, at the significantly lower temperatures and greater molecular fluxes of interstellar clouds, amorphous H_2O with diagnostic broad emission peaks (notably at 3.1 microns) is more likely to form.

Water that initially condenses in the amorphous form will not crystallize to cubic ice unless T rises above 110 K and it has sufficient time to anneal. However, 110 K is close to the sublimation temperature, while the sublimation rate of amorphous H_2O may be an order of magnitude greater than that of crystalline H_2O (Léger et al. 1983). Therefore, condensed H_2O phases may disappear (Kouchi et al. 1994). Indeed, one line of evidence for amorphous H_2O (as opposed to a crystalline form) in comets is the anomalously high sublimation rate of some comets (Mukai 1986).

Laboratory experiments on ice formation under low-pressure, cryogenic conditions suggest that the stable crystalline form of H_2O under circumstellar conditions is “cubic ice” (Gaffney and Matson 1980; Bartels-Rausch et al. 2012; Fuentes-Landete et al. 2015; Salzmann 2018). Note that fewer than 100 H_2O molecules are required to form a localized phase with the structural properties

of crystalline ice (Moberg et al. 2019; Jordan 2019). We recognize one likely circumstellar, low-temperature condensate, designated "*Circumstellar ice*."

Circumstellar ice: Occurring as < 0.05-micron-thick mantles on oxide and silicate dust grains, circumstellar cubic ice (H_2O) condenses in relatively cool ($50 < T < 150$ K) toroidal volumes surrounding some solar nebulas and O-rich stars (Omont et al. 1990).

Corundum (Al_2O_3): Corundum is thought to be the highest temperature oxide condensate at T ~ 1770 K (Grossman 1972; Davis and Richter 2014; Wood et al. 2019). However, because corundum reacts with the cooling gas to form other oxides such as hibonite or melilite it is a relatively rare mineral in CAIs.

CAI corundum: End-member Al_2O_3 with minor Mg and Fe is a rare mineral in CAIs. Sub-micron grains have been recorded in CAI cores, as inclusions in spinel or hibonite, or enclosing a central core of hibonite (Bar-Matthews et al. 1982; MacPherson et al. 1984; Wark 1986; Greshake et al. 1996).

URI corundum: Ma and Rossman (2009c) reported corundum associated with khamrabaevite and tistarite in an ultra-refractory inclusion from the Allende meteorite.

Tistarite (Ti_2O_3): The refractory mineral tistarite ($Ti^{3+}{}_2O_3$), a member of the corundum group, is known from a single 5- x 7-micron subhedral grain form a chondrule in the Allende

carbonaceous chondrite (Ma and Rossman 2009c; Ma et al. 2009a). It occurs in association with the primary phases corundum, khamrabaevite (TiC), and kaitianite ($Ti^{3+}_2Ti^{4+}O_5$), as well as rutile (TiO_2).

URI tistartite: Tistarite has been found as a single grain in an ultra-refractory inclusion contained within a chondrule from the Allende carbonaceous chondrite (Ma and Rossman 2009c).

Kaitianite ($Ti^{3+}_2Ti^{4+}O_5$): Two crystals of the mixed-valence titanium oxide mineral, kaitianite, were discovered by Ma (2019) in association with the primary phases corundum, tistarite, and khamrabaevite (TiC), as well as rutile (TiO_2). Micron-scale taitianite crystals have the monoclinic (*C*2/*c*) oxyvanite (V_3O_5) structure.

URI kaitianite: Kaitianite has been found as two grains in an ultra-refractory inclusion; the observed composition is $(Ti^{3+}_{1.75}Al_{0.05}Ti^{4+}_{0.10}Mg_{0.08}Fe_{0.02})Ti^{4+}O_5$ (Ma 2019).

Rutile (TiO_2): Rutile has been reported in association with tistarite and kaitianite by Ma and coworkers (Ma and Rossman 2009c; Ma et al. 2009a; Ma 2019), who suggest it is a primary phase in some URIs.

URI rutile: Occurs as micron-scale grains in an ultra-refractory inclusion from the Allende meteorite (Ma 2019).

Baddeleyite (ZrO_2): Presumably primary baddeleyite was described from meteorite MAC 88107, while grains of ZrO_2 (lacking structural details) have also been reported from Murchison and Allende chondrites (Krot et al. 2019). It occurs in association with allendeite and zirkelite.

URI baddeleyite: Occurs as micron-scale grains in URIs (Krot et al. 2019).

Anosovite [$(Ti^{4}, Ti^{3+}, Mg, Sc, Al)_3O_5$]: A second polymorph of Ti_3O_5 (in addition to kaitianite) was reported by Zhang et al. (2015) from an ultra-refractory inclusion in the Sayh al Uhaymir 290 (CH3) carbonaceous chondrite. This Sc-bearing phase has the orthorhombic (*Cmcm*) pseudobrookite structure and thus is equivalent to the discredited mineral "anosovite," which was originally identified in Ti-rich slags (Bowles 1988). It occurs as micron-scale grains in association with fassaite, spinel, anorthite, perovskite, panguite, davisite, and Fe-Ir-Mo-Os alloy nuggets.

URI anosovite: Two grains of ultra-refractory anosovite with the average composition of $(Ti^{4+}_{1.36}Ti^{3+}_{0.59}Mg_{0.34}Sc_{0.20}Al_{0.20}V_{0.05}Ca_{0.05}Si_{0.03}Fe_{0.03}Cr_{0.03}Zr_{0.03})O_5$ were reported by Zhang et al. (2015).

Spinel ($MgAl_2O_4$): Spinel is perhaps the most ubiquitous primary mineralogical component of CAIs in most chondrite types except CI (MacPherson 2014), as well as a common phase in Wark-Lovering rims. Fe-rich varieties of spinel group minerals, including chromite and hercynite, have

also been described from CAIs (e.g., Brearley and Jones 1998), but they are always of secondary origins.

CAI spinel: Most CAIs incorporate near end-member $MgAl_2O_4$ spinel, though with a wide range of observed trace and minor elements – notably Fe, Ti, V, Cr, and Zn (Brearley and Jones 1998). Spinel occurs in a variety of habits, including euhedral octahedral crystal inclusions in melilite, fassaite, or anorthite (Grossman 1980); as framboidal aggregates (El Goresy et al. 1979); surrounding a hibonite core (Steele 1995); as spinel-hibonite or spinel-perovskite spherules (Macdougall 1981); as spherical shells, or "palisades," of spinel enclosing melilite, fassaite, and/or anorthite, within larger CAIs (Wark and Lovering 1982b; Simon and Grossman 1997); in association with grossite-bearing inclusions (Brearley and Jones 1998); and as a common layer in Wark-Lovering rims (Wark and Lovering 1977).

AOA spinel: Spinel, typically in sub-micron grains, occurs as a primary phase associated with perovskite, fassaite, and anorthite in amoeboid olivine aggregates (Krot et al. 2004; Weisberg et al. 2004).

URI spinel: Spinel occurs in ultra-refractory inclusions in association with Sc-rich fassaite, REE-enriched perovskite, and other distinctive Sc, Zr, Ti, and REE phases (e.g., Ma and Rossman 2009b; Ma et al. 2013b, 2014a).

Hibonite ($CaAl_{12}O_{19}$): Hibonite, nominally $CaAl_{12}O_{19}$ but commonly incorporating significant $Mg+Ti^{4+} \leftrightarrow 2Al$ (up to ~4.5 and 9 wt. % MgO and TiO_2, respectively) as well as

minor V, Fe, Si, Cr, and Sc, is an important mineralogical component of CAIs in most types of chondrite meteorites (Keil and Fuchs 1971; Brearley and Jones 1998; MacPherson 2014). As much as a quarter of hibonite Ti may be present as Ti^{3+} - an indication of the highly reducing formation conditions of some CAIs (Ihinger and Stolper 1986; Beckett et al. 1988).

Hibonite is thought to form initially by reaction of the gas phase with corundum at T > 1700 K. Because of its high temperature of condensation (Davis and Richter 2014; Wood et al. 2019), hibonite is often the earliest preserved mineral in a CAI. As a consequence, hibonite petrology, major and trace element composition, and stable isotopes have been extensively studied to provide clues regarding primitive stellar environments (Hinton et al. 1988; Ireland 1988; Ireland et al. 1988; Brearley and Jones 1998; MacPherson 2014 and references therein; Kööp et al. 2016a, 2016b, 2018).

CAI hibonite: CAI hibonite occurs in a variety of contexts and morphologies (Ireland 1988), including acicular crystals associated with spinel, melilite, and perovskite in compact type A inclusions (Grossman 1975; Sylvester et al. 1993); as a common constituent of fluffy type A CAIs (Grossman 1975; Kornacki and Wood 1984); as 1- to 25-micron laths near the exterior of type B CAIs in the Allende CV chondrite (Blander and Fuchs 1975); in fine-grained spinel-rich inclusions (Kornicki and Wood 1985); as up to 1-millimeter diameter grains in CAI cores (Allen et al. 1980; Armstrong et al. 1982; MacPherson et al. 1983); in clusters of tabular crystals surrounded by spinel in a CAI core (Steele 1995); as inclusions in spinel, fassaite, melilite, and grossite CAI cores (Wark and Lovering 1977; Kornacki and Wood 1985; Mao et al. 1990; Weber and Bischoff 1994); in association with corundum (MacPherson et al. 1984; Hinton et al. 1988); as isolated crystal fragments (MacPherson et al. 1983; Ireland 1988); in spinel-hibonite spherules as 5- to 20-micron

bladed crystals (Macdougall 1981) and rims (MacPherson et al. 1984); and as layers in Wark-Lovering rims.

Perovskite ($CaTiO_3$): Near-stoichiometric calcium titanate perovskite, albeit with trace or minor Mg, Al, Si, Cr, Sc, V, Fe, Y, Zr, Nb, REE, Th, and U, is an important mineralogical component of CAIs in most chondrite types (Kornacki and Wood 1985; Weber and Bischoff 1994; Brearley and Jones 1998; MacPherson 2014). Twinning in perovskite crystallites suggest that they were heated above the cubic-orthorhombic transition at 1573 K.

CAI perovskite: Near end-member perovskite is a common phase in CAIs (Brearley and Jones 1998). Perovskite generally occurs as very fine-grained (< 15 microns) inclusions in melilite, spinel, and fassaite, and less commonly hibonite and grossite (Macdougall 1981; Fahey et al. 1994; Kojima et al. 1995; Steele 1995; Weber et al. 1995). It is also a common component in Wark-Lovering rims (Weisberg et al. 1993; Keller and Buseck 1994).

AOA perovskite: Perovskite, often as sub-micron grains in association with spinel, is a common primary phase in amoeboid olivine aggregates (Weisberg et al. 2004).

URI perovskite: Perovskite enriched in REE occurs in ultra-refractory inclusions, for example from the Allende carbonaceous chondrite in association with Sc-rich fassaite, spinel, and kangite (Ma and Rossman 2009b; Ma 2011; Ma et al. 2013b, 2014a, 2015).

Lakargiite ($CaZrO_3$): Ma (2011) identified sub-micron grains of a Zr-dominant perovskite in an ultra-refractory inclusion from the Acfer 094 carbonaceous chondrite. Lakargiite [$Ca_{0.95}(Zr_{0.87}Ti_{0.16})O_3$] is found as inclusions in hibonite and in association with perovskite [$Ca_{0.94}(Ti_{0.98}Zr_{0.06})O_3$], tazheranite, an Os-dominant alloy, and other refractory phases (Krot et al. 2019).

URI lakargiite: Ma (2011) identified Zr-rich perovskite from an ultra-refractory inclusion in Acfer 094.

Grossite ($CaAl_4O_7$): Grossite is one of several high-temperature calcium aluminate minerals that forms in CAIs, presumably in environments with low Mg and Si shortly after the condensation of corundum and hibonite (Michel-Lévy et al. 1982; Weber and Bischoff 1994; Aléon et al. 2002). Samples are near stoichiometric, though they commonly incorporate minor Mg, Si, Ti, and Fe.

CAI grossite: Grossite typically occurs as 5- to 10-micron diameter crystallites, often as inclusions in association with melilite, perovskite, spinel, hibonite, fassaite, and anorthite (Greenwood et al. 1992; Simon et al. 1994).

Krotite ($CaAl_2O_4$): Krotite is one of several refractory calcium aluminates that occur as primary condensates in CAIs with low Mg and Si (Ivanova et al. 2002; Ma et al. 2011b). Mikouchi

et al. (2009) identified a second polymorph of $CaAl_2O_4$, dmitryivanovite, which is thought to be a high-pressure form that results from impact transformation (see Part IV of this series).

CAI krotite: The single known occurrence of krotite was found as aggregates of crystals from 10- to 350-microns diameter in the central and mantle portions of a CAI from the NWA 1934 CV3 carbonaceous chondrite (Ivanova et al. 2002; Ma et al. 2011b). It is found in association with perovskite, melilite, grossite, hibonite, and spinel.

Machiite ($Al_2Ti_3O_9$): Krot et al. (2020) identified machiite, a new ultra-refractory oxide mineral from the Murchison carbonaceous chondrite that probably formed by direct condensation from a gas phase or by crystallization from a Ca-Al-rich melt in CAI-forming nebular regions. The ideal composition is $Al_2Ti_3O_9$ but, as is typical with ultra-refractory oxide phases, machiite incorporates significant Sc, Y, and Zr, as well. It was found as a single 4.4-micron diameter euhedral grain in association with euhedral corundum. Electron diffraction studies suggest a monoclinic (*C*2/*c*) schreyerite-type structure.

URI machiite: Krot et al. (2020) describe a single 4.4-micron diameter crystallite with empirical composition $(Al_{1.17}Sc_{0.56}Y_{0.10}Ti^{4+}_{0.08}Fe_{0.06}Ca_{0.03}Mg_{0.01})(Ti^{4+}_{2.71}Zr_{0.28}Si_{0.01})O_9$.

Zirkelite [$(Ti,Ca,Zr)O_{2-x}$]: Krot et al. (2019) record a possible grain of zirkelite, a complex Ti-Ca-Zr oxide with a defect cubic fluorite structure, as a micron-scale phase in an ultra-refractory inclusion from meteorite MAC 88107.

URI zirkelite: Zirkelite was identified as a micron-scale grain from MAC 88107 (Krot et al. 2019).

Kangite [$(Sc,Ti,Al,Zr,Mg,Ca)_{1.8}O_3$] and Panguite [$(Ti^{4+},Sc,Al,Mg,Zr,Ca)_{1.8}O_3$]: Ma and coworkers described two closely-related Ti-Sc oxides – Sc-dominant kangite (Ma et al. 2013b) and Ti-dominant panguite (Ma et al. 2012) from ultra-refractory inclusions in the Allende carbonaceous chondrite, as well as an occurrence of panguite from the Murchison chondrite (Ma et al. 2011a). Both minerals occur as micron-scale grains in association with Sc-rich davisite, perovskite, and spinel, and both have cation-deficient bixbyite structures. Thus, kangite and panguite might be lumped into a single natural kind; however, kangite is reported to be cubic *Ia*{bar3}, whereas the panguite structure occurs in the orthorhombic subgroup *Pbca*. Therefore, we list two different natural kinds, subject to further structural and compositional details. Note that Ma et al. (2012) also report several Zr-rich grains of panguite.

URI kangite: Sc-dominant kangite was described from an ultra-refractory inclusion in the Allende carbonaceous chondrite (Ma et al. 2013b).

URI panguite: Micron-scale crystallites of Ti-dominant panguite are associated with fassaite in ultra-refractory inclusions (Ma et al. 2012; Krot et al. 2019). The most complete description is from a 20- x 30-micron amoeboid olivine aggregate in the Allende carbonaceous chondrite, while additional occurrences have been reported from the Allende, Murchison, and Sayh al Uhaymir 290 meteorites (Ma et al. 2012; Zhang et al. 2015).

Zirconolite ($CaZrTi_2O_7$): Ma and Rossman (2008) reported zirconolite in an ultra-refractory inclusion in an amoeboid olivine aggregate from the Allende carbonaceous chondrite. It is associated with micron-scale inclusions of cubic zirconia (tazheranite) and Fe-Ni and PGE alloys.

URI zirconolite: Zirconolite from the Allende meteorite was recorded by Ma and Rossman (2008).

Tazheranite [$(Zr,Sc,Ca,Y,Ti)O_{1.75}$]: Ma and Rossman (2008) described natural Sc-Ti-rich cubic zirconia as crystals up to 1.2-micron diameter from ultra-refractory inclusions within an AOA in the Allende meteorite. Tazheranite occurs as inclusions in zirconolite; associated minerals include fassaite and Fe-Ni and PGE alloys. A second occurrence from an Allende URI in fine-grained matrix (Ma et al. 2014a) is associated with allendeite, spinel, fassaite, and perovskite.

URI tazheranite: Tazheranite is a rare mineral from ultra-refractory inclusions from the Allende meteorite (Ma and Rossman 2008; Ma et al. 2014a; Krot et al. 2019).

Allendeite ($Sc_4Zr_3O_{12}$): Allendeite was described by Ma et al. (2014a), who examined grains up to 25-microns diameter in an ultra-refractory inclusion in the Allende carbonaceous carbonate. It contains inclusions of spinel and hexamolybdenum, and is closely associated with fassaite, perovskite, tazheranite, and Os-dominant PGE alloys.

URI allendeite: Allendeite was discovered in an ultra-refractory inclusion from the Allende meteorite (Ma et al. 2014a; Krot et al. 2019).

SILICATES

Refractory silicates are major components of CAIs, AOAs, and URIs. Several major rock-forming mineral groups, including olivine, garnet, pyroxene, and feldspar, are represented, as well as rare ultra-refractory Ti, Sc, and Zr silicates.

Quartz (SiO_2): Komatsu et al. (2018) reported an unusual occurrence of silica condensation within the solar nebula in a region with low Mg/Si – an example of nebular element fractionation in the period following CAI formation. Quartz occurs as grains up to 20-microns diameter in association with fassaite, anorthite, and spinel in an amoeboid olivine aggregate from the Yamato-793261 carbonaceous chondrite.

AOA quartz: Quartz occurs as primary grains associated with fassaite, forsterite, anorthite, and spinel in an AOA (Komatsu et al. 2018).

Olivine group [$(Mg,Fe,Ca,Mn)_2SiO_4$]: Nebular olivine represents a solid solution among four principal end members, forsterite (Mg), fayalite (Fe), monticellite (Ca), and tephroite (Mn). Primary olivine that formed by direct condensation or melt solidification in CAIs and AOAs is typically close to end-member forsterite (Brearley and Jones 1998). Davis et al. (1991) describe rare instances where CAI forsterite (replaced by akermanite) and fassaite (replaced by gehlenite and perovskite) have been partially evaporated, resulting in a different mode of mineral formation

1249 due to a reheating event (and loss of Si and Mg) to at least 1700 K, possibly by remixing close to
1250 the Sun (Davis and Richter 2014).
1251
1252 *CAI forsterite:* Forsteritic olivine is a minor primary phase in CAIs. It occurs, for example, as
1253 < 5-micron diameter grains associated with spinel and fassaite, as well as in Wark-Lovering rims,
1254 in a variety of carbonaceous chondrites (Greenwood et al. 1994; MacPherson and Davis 1994),
1255 and as an interstitial phase with anorthite and melilite in a CAI from PCA91082 (Birjukov and
1256 Ulyanov 1996).
1257
1258 *AOA forsterite:* Mg-rich olivine (Fo > 98), in some cases with rims enriched in Mn (Mn > Fe),
1259 is the principal constituent of AOA's (Krot et al. 2004; Weisberg et al. 2004). Micron-scale crystals
1260 (<10 microns) occur in irregular-shaped aggregates up to 1 millimeter in diameter. Klöck et al.
1261 (1989) suggest that the Mn enrichment arose from condensation of a tephroite component directly
1262 from the solar nebular at temperatures close to 1100 K, compared to the ~1440 K condensation of
1263 pure forsterite, and far above the ~500 K reaction of Fe metal with forsterite to form fayalite.
1264
1265 **Garnet group [Ca_3(Al, Ti^{3+},V,Sc)$_2Si_3O_{12}$]:** Ca-Al garnet with significant Ti^{3+} and/or Sc is a
1266 scarce primary phase in the earliest nebular condensates. Garnet with composition close to the
1267 grossular end-member (ideally $Ca_3Al_2Si_3O_{12}$) is known almost exclusively as a secondary phase
1268 in chondrites (Fuchs 1974; Wark et al. 1987; Brearley and Jones 1998; Rubin and Ma 2017). In
1269 addition, Simon and Grossman (1992) reported goldmanite [Ca_3(V,Al,Fe,Ti)$_2Si_3O_{12}$] in a
1270 Fremdlinge from the Leoville chondrite (see Part III). However, Ma and coworkers have

discovered two rare primary garnets with the general formula $Ca_3(Ti^{3+},Sc,Mg,Y)Si_3O_{12}$ in chondrite meteorites (Ma 2012; Ma et al. 2017a), including the end-members eringaite ($Ca_3Sc_2Si_3O_{12}$) and rubinite ($Ca_3Ti^{3+}_2Si_3O_{12}$). These compositional extremes reveal fascinating heterogeneities in the early solar nebula.

Eringaite (Ma 2012), the Sc-dominant, Al-poor garnet $Ca_3(Sc,Y,Ti)_2Si_3O_{12}$, was discovered with rubinite in an ultra-refractory inclusion within an AOA from the Vigarano carbonaceous chondrite. It occurs as micron-scale crystals within davisite and in association with spinel, tazheranite, and hexaferrum.

The Ti^{3+} end-member, dubbed rubinite by Ma et al. (2017a), was discovered as crystals up to 20-microns diameter in both type A CAIs from the Allende and Efremovka carbonaceous chondrites, where it occurs with the major primary CAI minerals, as well as eringaite and Ti^{3+}-dominant fassaite (grossmanite). Rubinite significantly enriched in Y, Sc, and Zr was found in an ultra-refractory inclusion in the Vigarano meteorite with spinel, panguite, fassaite, and davisite, which is enclosed in an AOA.

CAI rubinite: Rubinite occurs in type A CAIs from the Allende and Efremovka carbonaceous chondrites (Ma 2012).

URI rubinite: Rubinite occurs with eringaite in an ultra-refractory inclusion in the Vigarano chondrite (Ma et al. 2017a); significantly enriched in Y, Sc, and Zr.

URI eringaite: Eringaite occurs with rubinite in an ultra-refractory inclusion in the Vigarano chondrite (Ma et al. 2017a).

Melilite group [gehlenite ($Ca_2Al_2SiO_7$) to åkermanite ($Ca_2MgSi_2O_7$)]: The calcium silicates of the melilite group, which feature complete solid solution between Al_2 (gehlenite) and MgSi (åkermanite) end members, are major mineralogical constituents of CAIs. Unlike many igneous examples, CAI melilite contains minimal Fe and Na. Wood et al. (2019), in revised calculations of condensation temperatures that take trace elements into account, suggest that gehlenite-dominant melilite condenses first, between 1550 and 1600 K, following the appearance of corundum and hibonite. As condensation temperatures fall, melilite crystals typically develop zoning from gehlenite-rich cores to åkermanite-dominant rims. Because melilite in CAIs spans the complete compositional range from Al_2 to MgSi end-members, we lump CAI gehlenite and åkermanite into a single natural kind. Melilite is common in most types of CAIs, as well as in AOAs.

CAI melilite: Melilite is among the commonest primary condensed phases in CAIs, ranging in composition from $Åk_{01}$ to $Åk_{100}$ (Grossman 1975, 1980; Wark et al. 1987; Davis et al. 1991; Podosek et al. 1991), often with zoning from more gehlenite-rich cores to more åkermanite-rich rims. It occurs as coarse-grained (2 to 3 millimeter) crystals in the cores of many CAIs (e.g., MacPherson 2014); as rims on spinel (Holmberg and Hashimoto 1992); in nodular aggregates with spinel and fassaite (Weisberg et al. 1993); as a component of hibonite-, spinel-, and grossite-

bearing inclusions (J.N. Grossman et al. 1988; Weber et al. 1995); and as layers in Wark-Lovering rims (Wark and Lovering 1977).

AOA melilite: Gehlenite-rich melilite (Ge_{80-85}) is an occasional primary phase in AOAs, typically associated with anorthite and spinel (Krot et al. 2004). Weisberg et al. (2004) suggest that some AOA melilite underwent solid-state reactions on cooling to form anorthite.

URI melilite: Ma et al. (2015) reported gehlenitic melilite in a Sc-rich ultra-refractory inclusion in the Vigarano carbonaceous chondrite, associated with warkite, davisite, perovskite, and spinel.

Clinopyroxene Group [$Ca(Mg,Al,Ti^{3+},Ti^{4+},Sc,V)(Al,Si)SiO_6$]: Fassaite and related calcic clinopyroxenes are abundant primary phases in CAIs, AOAs, and URIs (Brearley and Jones 1998; Davis and Richter 2014; MacPherson 2014; Sack and Ghiorso 2017). Ca-Mg-dominant, Fe-poor clinopyroxene with significant Al and Ti (both Ti^{3+} and Ti^{4+}), and occasionally Sc or V, is closest compositionally to the IMA-approved species diopside ($CaMgSi_2O_6$), though end-member diopside has not been confirmed as a primary nebular condensate (Brearley and Jones 1998). The discredited name "fassaite" (Morimoto et al. 1988) is still widely used by the meteoritics community and the name is retained here.

The crystal chemistry of fassaite can be modeled with four dominant ideal end-members (Sack and Ghiorso 1994a, 1994b, 1994c, 2017). In addition to diopside, rare Al-rich grains are closer to ($CaAl_2SiO_6$), the IMA-approved end-member kushiroite (Kimura et al. 2009), though long called the "calcium Tschermak's pyroxene" (e.g., Ma et al. 2009b). Other examples are closer in

composition to grossmanite ($CaTi^{3+}AlSiO_6$; Ma and Rossman 2009a), while Sack and Ghiorso (1994a, 1994c, 2017) also recognize "alumino-buffonite" ($CaMg_{0.5}Ti^{4+}_{0.5}AlSiO_6$) as an important additional end-member with the maximum possible octahedral Ti^{4+} content. Phase equilibria in this four-component fassaite system are complex; Sack and Ghiorso (2017) suggest as many as four miscibility gaps, based on both thermodynamic modeling and observed coexistence of fassaite rims and cores of different compositions that appear to be in equilibrium (Wark and Lovering 1982). Additional complexity arises from the occurrence of rare micron-scale grains of burnettite ($CaV^{3+}AlSiO_6$; Ma and Beckett 2016) and davisite ($CaSc^{3+}AlSiO_6$; Simon et al. 1996; Ma and Rossman 2009b).

How many natural kinds of primary nebular clinopyroxene should be recognized remains an open question. Evidence for fassaite miscibility gaps points to the need for multiple natural kinds; however, until more analyses and other diagnostic attributes are available for cluster analysis, we lump primary nebular clinopyroxenes from the quadrilateral defined by diopside, kushiroite, grossmanite, and "alumino-buffonite" into one natural kind, "fassaite" [$Ca(Mg,Al,Ti^{3+},Ti^{4+})(Al,Si)SiO_6$].

The fassaite phase region may extend to Sc- and/or V-rich compositions; however, only two examples of these extremes have been reported and both appear idiosyncratic. Burnettite is an oddity, with octahedral M2 site composition [$(V^{3+}_{0.29}Sc_{0.24}Ti^{3+}_{0.13}Al_{0.09})Ti^{4+}_{0.12}Mg_{0.08}$] (Ma and Beckett 2016); thus, V, Sc, and Ti are present in roughly equal amounts and no cation is present at greater than 29 mol % M2 occupancy. Burnettite is therefore far from any ideal end-member

composition. We designate the only known occurrence, from an unusual V-rich fluffy type A CAI, as "*CAI burnettite*."

Davisite [$Ca(Sc,Ti^{3+},Ti^{4+},Mg,Zr)AlSiO_6$], is a minor phase that occurs occasionally in ultra-refractory inclusions (Ma and Rossman 2009b). Common associations include fassaite (with significantly less Sc than in coexisting davisite), as well as perovskite, spinel, and a host of rare Sc-bearing minerals, including eringaite, kangite, panguite, rubinite, thortveitite, and warkite (Ma et al. 2011a, 2012, 2013b, 2015, 2017a; Ma 2012; Krot et al. 2019).

CAI fassaite: Ca-rich, Fe-poor pyroxenes are, in combination with spinel, the most common CAI phases, occurring in types A, B, and C CAIs and URIs (Grossman 1980; Wark 1987; Wark et al. 1987; Podosek et al. 1991; Brearley and Jones 1998), as well as layers in Wark-Lovering rings (Wark and Lovering 1977). Primary fassaite commonly occurs with anorthite, melilite, forsterite, and spinel as CAI cores, mantles, and rims (Macdougall 1979, 1981; Doukhan et al. 1991; Kimura et al. 1993; Kojima et al. 1995; Simon et al. 1996). Of special note regarding CAI fassaite is the occurrence of both trivalent and tetravalent Ti—a consequence of the extremely reducing conditions of its formation (Beckett 1986), perhaps dominated by hot H_2 gas, as well as conditions in which $C/O > 0.5$.

AOA fassaite: Al-Ti-rich calcic clinopyroxene with Al and Ti as high as 20 and 13 wt. %, respectively, is a common primary phase in AOAs, especially in Al-rich refractory inclusions where it occurs in close association with forsterite and anorthite, as well as melilite, spinel, and perovskite (Hashimoto and Grossman 1987; Krot et al. 2004; Weisberg et al. 2004; Ma et al. 2012).

URI fassaite: Fassaite enriched in Sc, Ti^{3+}, and/or V^{3+} occurs as sub-millimeter grains in ultra-refractory inclusions in association with davisite, spinel, perovskite, melilite, and a variety of rare oxides and silicates (Ma and Rossman 2009a, 2009b; Ma et al. 2013b, 2014a, 2015).

CAI burnettite: V-dominant, Sc- and Ti-rich calcic clinopyroxene was identified by Ma and Beckett (2016) from one fluffy type A CAI in the Allende meteorite.

URI davisite: Davisite is a Sc-dominant calcic clinopyroxene that occurs in association with fassaite, perovskite, spinel, and rare Sc-bearing phases in ultra-refractory inclusions (Ma and Rossman 2009b).

Feldspar Group [$(Na,Ca)(Al,Si)_4O_8$]

Anorthite ($Ca_2Al_2SiO_8$): End-member (i.e., Na-free) anorthite is common as both a primary and secondary mineral, notably in type B and type C CAIs (Podosek et al. 1991; Caillet et al. 1993; Brearley and Jones 1998). Condensing initially at ~1410 K (Davis and Richter 2014), anorthite is the last of the common CAI primary phases to appear; it crystallizes from a melt and is often found in close association with melilite or spinel (MacPherson and Davis 1993; Kojima et al. 1995). Note that plagioclase feldspar, typically with a significant albitic component ($NaAlSi_3O_8$), is a common secondary phase in chondrite meteorites, including as a fine-grained alteration phase in CAIs and their Wark-Lovering rims, as a common minor phase in AOAs, and in chondrules (see Part III).

1399 *CAI anorthite:* The last major primary phase to condense from the solar nebula, anorthite is
1400 commonly found as coarse-grained laths in CAIs (Brearley and Jones 1998).
1401
1402 *AOA anorthite:* Near end-member Ca plagioclase is a common, if volumetrically minor,
1403 constituent of AOAs (Weisberg et al. 2004).
1404
1405 **Dmisteinbergite ($CaAl_2Si_2O_8$):** Rare individual crystals of dmisteinbergite, a hexagonal high-
1406 temperature polymorph of anorthite, have been documented from a type B FUN CAI in the Allende
1407 carbonaceous chondrite (Ma et al. 2013a). This phase points to an origin environment with T >
1408 1500 K and P < 10^{-6} atm (Abe et al. 1991; Mendybaev et al. 2009), likely close to the protosun,
1409 where it crystallized directly from a silicate vapor or melt phase. Ma et al. (2013a) also note a Ba-
1410 rich grain of dmisteinbergite with up to 27 atom percent Ba substituting for Ca.
1411
1412 *CAI dmisteinbergite:* Dmisteinbergite occurs as 100- to 600-micron diameter crystals in
1413 association with melilite, fassaite, and spinel (Ma et al. 2013a).
1414
1415 **Baghdadite [$Ca_3(Zr,Ti)Si_2O_9$]:** Ma (2018) reported the first extraterrestrial occurrence of
1416 baghdadite, a CAI silicate mineral from the Allende carbonaceous chondrite that may be one of
1417 the earliest primary silicate condensates. A single 0.8-micron diameter euhedral grain is associated
1418 with primary spinel, perovskite, hibonite, fassaite, refractory metal nuggets, and the rare phases
1419 burnettite and paqueite as inclusions in melilite. The baghdadite structure is monoclinic ($P2_1/a$).
1420

CAI baghdadite: Ma (2018) reported a single baghdadite grain with composition $(Ca_{2.77}Mg_{0.08})(Zr_{0.55}Ti_{0.35}Nb_{0.02})(Si_{1.89}Al_{0.35})O_9$.

Rhönite [$Ca_2(Mg,Al,Ti)_6(Si,Al)_6O_{20}$], Addibischoffite [$Ca_2(Al,Mg,V,Ti)_6(Al,Si)_6O_{20}$), and Warkite [$Ca_2(Sc,Ti,Al,Mg,Zr)_6Al_6O_{20}$]: Rhönite, a rare refractory member of the sapphirine group (triclinic, *P*{bar1}), was first reported by Fuchs (1971), and was subsequently identified in both type A and type B CAIs in association with melilite, spinel, fassaite, and perovskite (Fuchs 1978; Grossman 1980; Podosek et al. 1991; Jambon and Boudouma 2011). Ma and colleagues subsequently identified two closely-related Ca-Al oxides/silicates, also with the *P*{bar1} sapphirine structure. Addibischoffite, with the general formula $[Ca_2(Al,Mg,V,Ti)_6(Al,Si)_6O_{20}]$, was discovered as a 9-micron diameter crystal in a CAI from the Acfer 214 carbonaceous chondrite. Like rhönite, it occurs in association with typical CAI phases – Al-dominant fassaite (kushiroite), hibonite, spinel, melilite, perovskite, anorthite, and Fe-Ni alloy (Ma et al. 2017b). We conclude that rhönite and addibischoffite are part of a continuous solid solution and form in similar environments. Therefore, we combine these minerals into one natural kind: "*CAI rhönite*."

Warkite was identified by Ma et al. (2015) from ultra-refractory inclusions in the Murchison and Vigarano carbonaceous chondrites. Warkite, general formula $[Ca_2(Sc,Ti,Al,Mg,Zr)_6Al_6O_{20}]$, occurs as aggregates of crystals up to 4-microns diameter in association with perovskite, davisite, spinel (in Murchison), and melilite (in Vigarano). Warkite may well form a continuous solid-solution with rhönite and addibischoffite; however, because it forms in a different mineralogical environment and appears to be Si-poor, we recognize "URI warkite" as a distinct natural kind.

CAI rhönite: Rhönite has been identified in both A and B type CAIs from several carbonaceous chondrites (Fuchs 1971; Ma et al. 2017b).

URI warkite: Warkite was described by Ma et al. (2015) from ultra-refractory inclusions in the Murchison and Vigarano meteorites (see also, Krot et al. 2019).

Paqueite [$Ca_3TiSi_2(Al,Ti,Si)_3O_{14}$]: Ma and Beckett (2016) identified a new Ca-Ti silicate from a fluffy type A CAI in the Allende carbonaceous chondrite. It occurs as micron-scale euhedral crystals in association with melilite, spinel, V-rich perovskite, fassaite, hibonite, and refractory metal grains.

CAI paqueite: Paqueite was discovered by Ma and Beckett (2016) in a V-rich fluffy type A CAI.

Thortveitite ($Sc_2Si_2O_7$): Ma et al. (2011a) and Ma (2012) report the occurrence of thortveitite in an ultra-refractory inclusion from the Murchison meteorite. It occurs with fassaite, davisite, panguite, and spinel as subhedral crystals up to 9-microns maximum dimension.

URI thortveitite: Thortveitite occurs in an ultra-refractory inclusion in the Murchison chondrite (Ma et al. 2011a; Krot et al. 2019).

Silicate Glass (Ca,Mg,Al,Si,O): A glass of aluminous pyroxene composition is a significant component of hibonite-silicate spherules, a type of CAI less than 200-microns in diameter (Kimura

et al. 1993; Beckett and Stolper 1994; Russell et al. 1998; MacPherson 2014). Silicate glass occurs with fassaite and hibonite, often in association with perovskite, melilite, and grossite. These spherules contain primary phases that crystallized from a melt or, in the case of hibonite, may represent unmelted relict crystals. Note that we distinguish nebular silicate glass, which formed from a melt, from stellar amorphous silicate, which condenses from the gas phase (Hazen and Morrison 2020).

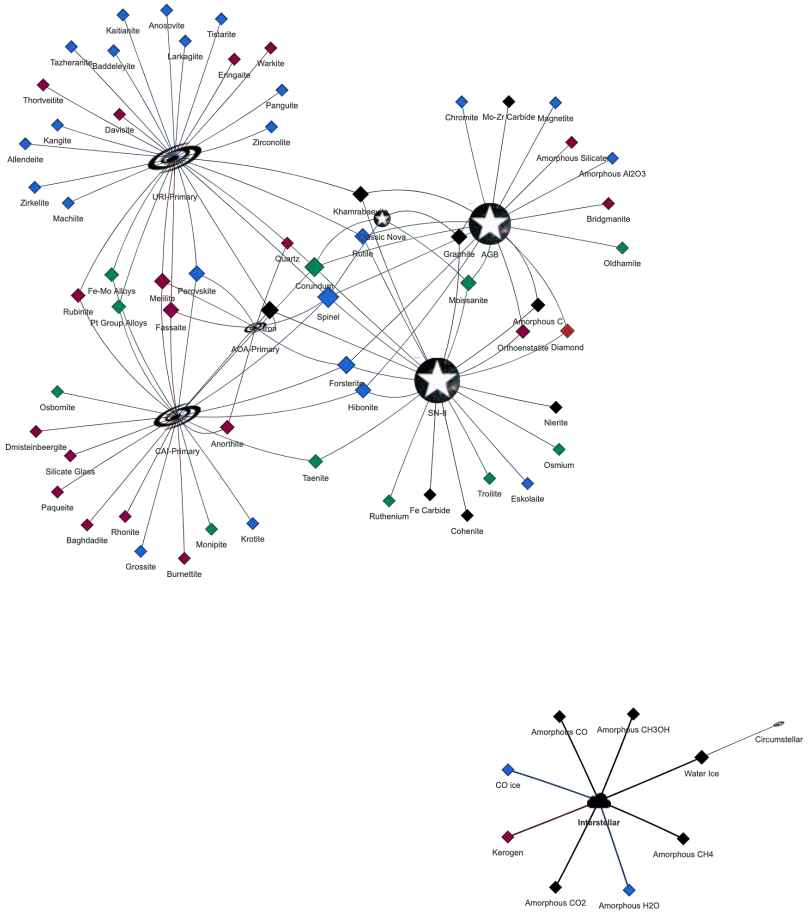
CAI silicate glass: Silicate glass is a significant component of hibonite-silicate spherules (MacPherson 2014).

NETWORK GRAPH OF STELLAR AND PRIMARY NEBULAR MINERALS

The evolutionary system of mineralogy is illustrated using bipartite mineral network graphs (Fruchterman and Reingold 1991; Asratian et al. 1998; Morrison et al. 2017), which display relationships among mineral phases and their attributes, in this instance their paragenetic modes (Hazen et al. 2019; Hazen and Morrison 2020; Morrison et al. 2020). Figure 4 displays a bipartite force-directed network graph of primary stellar, interstellar, and nebular minerals formed prior to ~4,565 Ma, in which 69 different phases, including 10 amorphous condensed phases, are represented by diamond-shaped nodes. Each of these mineral nodes is linked to one or more node representing a paragenetic mode of formation. Three different star-shaped nodes (AGB, SN-II, and CNova) represent stellar environments that impart distinctive isotopic signatures to minerals. A cloud-shaped node indicates interstellar dense molecular clouds (DMC), whereas four flattened disk icons represent different primary mineral-forming nebular environments (Circumstellar, CAI, AOA, and URI).

Node size, shape, and color convey information. Mineral compositions are indicated by the color of diamond-shaped mineral nodes: black (C-bearing), green (lacking C or O), blue (contains O, but not C or Si), and red (contains Si + O). The sizes of mineral nodes correspond to the numbers of paragenetic modes to which they are linked. Similarly, the sizes of the star-, cloud-, and disk-shaped symbols indicate the numbers of different minerals to which they are associated.

At this early stage of mineral evolution, 8 different low-temperature interstellar and nebular condensed molecular phases ($T < 100$ K) form a separate network from 56 high-temperature stellar and nebular condensates ($T > 1100$ K). In future parts of this series, phases formed at intermediate temperatures in planetary surface environments will provide links between these two mineral-forming environments.



1498

1499 Figure 4. Bipartite force-directed network graph (Morrison et al. 2017) of primary stellar, interstellar, and
1500 nebular minerals linked to their modes of paragenesis. Diamond-shaped nodes represent condensed
1501 crystalline and amorphous phases [black (C-bearing), green (not C or O), blue (contains O, but not C or Si),
1502 and red (contains Si + O)]. Star-shaped nodes represent three types of host stars—asymptotic giant branch
1503 stars (AGB), Type II supernovae (SN-II), and classical novae (CNova); the cloud-shaped node represents
1504 dense molecular clouds (DMC); and four disk-shaped nodes indicate circumstellar environments, CAIs,
1505 AOAs, and URIs. The sizes of nodes correspond to the numbers of links to other nodes. Note that 8 low-
1506 temperature phases of the interstellar medium are not linked to 61 high-temperature primary phases of
1507 stellar and nebular environments.

1508

1509 This bipartite network of mineral evolution is a visual representation of all confirmed stellar,
1510 interstellar, and primary nebular minerals described in Parts I and II of the evolutionary system of
1511 mineralogy. As new parts are introduced, and new nodes for minerals and paragenetic processes

1512 are added, this information-rich graphical approach will provide a dynamic, expanding, interactive
1513 view of the entire sweep of mineral evolution.
1514

IMPLICATIONS

Stage II of mineral evolution introduces for the first time several mineral groups that played important roles in planetary evolution, including garnet, melilite, clinopyroxene, and quartz. Like several of the phases described among the stellar condensation minerals of Stage I, many of which occur in 2 or 3 different natural kinds based on very different isotopic signatures associated with different kinds of stars (Hazen and Morrison 2020), several primary nebular minerals occur as condensates in CAIs, AOAs, and URIs – different mineral-forming environments that impart distinctive combinations of compositional, morphological, and petrologic attributes.

As with Part I of this series, we emphasize that this contribution is a preliminary foray into a rich and rapidly evolving topic of mineralogical investigation. New nebular minerals, notably micron-scale refractory phases, are being discovered every year; unusual varieties of chondrite meteorites continue to be discovered; and new and more precise analytical measurements of mineral ages, isotopic idiosyncrasies, and nanoscale chemical and textural features refine our understanding of this earliest stage of solar system evolution. We also note that condensed non-crystalline phases play an important role in interstellar and nebular mineralogy, as they do in other episodes of Earth's mineral evolution. Further efforts to characterize glasses and amorphous condensates will undoubtedly reveal new natural kinds, as they clarify our understanding of heating and cooling events in the dynamic, heterogeneous nebular environment.

This contribution, the second in a series that considers chronologically the evolutionary sequence of mineral formation, will be followed by an examination of the primary igneous minerals of chondrules, which are the most abundant constituents of chondrite meteorites (Part III). The primary minerals of differentiated asteroidal bodies, as well as numerous secondary minerals that arise through thermal, aqueous, and impact alteration of chondrite minerals in

planetesimal environments – minerals as preserved in altered chondrite and achondrite meteorites – will provide the focus of Parts IV and V.

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Table 2. Diagnostic properties of circumstellar and interstellar condensed phases. Unconfirmed phases appear in [brackets].

Group	Species (Formula)	Natural Kind	Characteristics	References
NATIVE ELEMENTS				
	Hydrogen (H_2)	[Interstellar hydrogen]	A fraction of H_2 may condense at T ~ 10 K	1
	Nitrogen (N_2)	[Interstellar nitrogen]	Molecular nitrogen should condense heterogeneously at T < 20 K	2,3
	Oxygen (O_2)	[Interstellar oxygen]	Most molecular oxygen reacts with H_2 to form H_2O	4,5
OXIDES				
	Water (H_2O)	Interstellar cubic ice	Diagnostic sharp IR emission features at 44 and 60 microns	6-8
		Interstellar amorphous H_2O	Diagnostic O-H stretch at 3.05 microns	7-10
	Carbon Monoxide (CO)	Interstellar amorphous CO	Diagnostic absorption at 4.67 microns	11,12
		Interstellar CO	Forms when amorphous CO anneals at T > 23 K	13
	Carbon Dioxide (CO_2)	Interstellar amorphous CO_2	IR absorption features at 4.27 and 15.2 microns	14
	Sulfur Dioxide (SO_2)	[Interstellar SO_2]	Diagnostic absorption at ~7.6 microns	15,16
ORGANIC MOLECULAR SOLIDS				
	Methanol (CH_3OH)	Interstellar amorphous CH_3OH	Diagnostic absorption at 3.54, 3.95, 8.9, and 9.75 microns	8,17
	Methane (CH_4)	Interstellar amorphous CH_4	Diagnostic 7.676-micron absorption	8,18

Cyanide (XCN; X = H,O)	[Interstellar *X*CN]	Diagnostic absorption at 4.62 microns	8,19,20
Formaldehyde (H_2CO)	[Interstellar H_2CO]	5.81 and 5.83-micron absorption; 4830 MHz emission	8,21
Formic Acid (HCOOH)	[Interstellar HCOOH]	Diagnostic absorption at 5.85 and 7.243 microns	8,22,23
Acetaldehyde (CH_3HCO)	[Interstellar CH_3HCO]	Diagnostic 7.414-micron absorption	8,23
Carbonyl Sulfide (OCS)	[Interstellar OCS]	Diagnostic 4.91-micron absorption	8,24,25
Ammonia (NH_3)	[Interstellar NH_3]	Suggested by bands at 3.5 and ~9 microns	8,26
Kerogen (C,H,N,O)	Interstellar kerogen	Preserved in relatively unaltered carbonaceous chondrites; Diagnostic absorption at 3.3, 3.47, 6.2, 8.6, and 11.3 microns	8,27-31

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Table 3. Properties of primary phases in the solar nebula formed by condensation, melt crystallization, and solid-state reactions.

***CAI* = calcium-aluminum-rich inclusion; *AOA* = amoeboid olivine aggregate; *URI* = ultra-refractory inclusion.**

Group	Species (Formula)	Natural Kind	Characteristics	References
			NATIVE ELEMENTS	
	Pt group alloys (Pt,Ru,Os, etc.)	*CAI PGE alloy*	Nano- to micro-scale alloys, space group $P6_3/mmc$.	1,2
		URI PGE alloys	Often Os-dominant	3-5
	Iron (Fe,Ni) [also called "kamacite"]	*CAI iron*	Occurs as a primary phase in type A, B, and C CAIs	1,6-8
		AOA iron	Ubiquitous in AOAs, typically with 5 to 7 wt % Ni	9-11
		URI iron	Sub-micron grains associated with ultra-refractory minerals	6
	Taenite (Fe,Ni)	*CAI taenite*	Occurs in type A and B CAIs; typically 10 to 50 wt. % Ni	1,6,8
	Fe-Mo alloys (Fe,Mo,Ru,Os,etc.)	*CAI Fe-Mo alloys*	Includes IMA-approved hexaferrum and hexamolybdenum	12-14
		URI Fe-Mo alloys	Includes IMA-approved hexaferrum and hexamolybdenum	4,5,14
			CARBIDES	
	Khamrabaevite (TiC)	*URI khamrabaevite*	Associated with corundum and tistarite	15
			NITRIDES	
	Osbornite (TiN)	*CAI osbornite*	As a rare phase in CAIs	16,17
			PHOSPHIDES	
	Monipite (MoNiP)	*CAI monipite*	Known from micron-scale grains in the Allende meteorite	18
			OXIDES	
	Ice (H_2O)	*Nebular cubic ice*	Condenses at $50 < T < 150$ K in the outer nebula	19

Corundum (Al_2O_3)	*CAI corundum*	Near end-member; in CAI cores, inclusions in hibonite or spinel	1,20,21
	URI corundum	Associated with khrambaevite and tistarite	15
Tistarite (Ti_2O_3)	*URI tistarite*	Associated with khrambaevite and corundum	15,22
Kaitianite ($Ti^{3+}_2Ti^{4+}O_3$)	*URI kaitianite*	Known from two grains, both micron-scale	15,22
Rutile (TiO_2)	*URI rutile*	Associated with tistarite and kaitianite	15,22
Baddeleyite (ZrO_2)	*URI baddeleyite*	Associated with zirkelite	23
Anosovite [$(Ti^4,Ti^{3+},Mg,Sc,Al)_3O_5$]	*URI anosovite*	Known from two grains, both micron-scale	14
Spinel ($MgAl_2O_4$)	*CAI spinel*	Ubiquitous in CAIs; a component of Wark-Lovering rims	1,2
	AOA spinel	Associated with perovskite, fassaite, and anorthite	10,11
	URI spinel	Associated with Sc-rich fassaite, REE-enriched perovskite	5,24,25
Hibonite ($CaAl_{12}O_{19}$)	*CAI hibonite*	Common in CAIs and Wark-Lovering rims	1,2
Perovskite ($CaTiO_3$)	*CAI perovskite*	Common in CAIs and Wark-Lovering rims	1,2
	AOA perovskite	Common as sub-micron grains associated with spinel	1,11
	URI perovskite	Typically enriched in REE; associated with Sc minerals	4,5,24-26
Lakargiite ($CaZrO_3$)	*URI lakargiite*	Sub-micron grains as inclusions in hibonite	4
Grossite ($Ca_2Al_4O_7$)	*CAI grossite*	5- to 10-micron grains as inclusions in major CAI phases	27,28
Krotite ($CaAl_2O_4$)	*CAI krotite*	One occurrence; in association with major CAI phases	29,30
Machiite ($Al_2Ti_3O_9$)	*URI machiite*	Known from a single 4.4-micron diameter crystallite	31

Zirkelite [$(Ti,Ca,Zr)O_{2-x}$]	*URI zirkelite*	Found in association with baddeleyite	23
Kangite [$(Sc,Ti,Al,Zr,Mg,Ca)_{1.8}O_3$]	*URI kangite*	Cation-deficient, cubic bixbyite structure	25
Panguite [$(Ti,Sc,Al,Mg,Zr,Ca)_{1.8}O_3$]	*URI panguite*	Cation-deficient, orthorhombic bixbyite structure	12,14,32,33
Zirconolite ($CaZrTi_2O_7$)	*URI zirconolite*	Associated with tazheranite and metal alloys	3
Tazheranite [$(Zr,Sc,Ca,Y,Ti)O_{1.75}$]	*URI tazheranite*	Associated with zirconolite and metal alloys	3,13
Allendeite ($Sc_4Zr_3O_{12}$)	*URI allendeite*	Contains spinel and refractory metal inclusions	5

SILICATES

Quartz (SiO_2)	*AOA quartz*	Associated with fassaite, forsterite, anorthite, and spinel	34
Olivine Group [$(Mg,Fe,Ca,Mn)_2SiO_4$]			
Forsterite Mg_2SiO_4	*CAI forsterite*	A minor primary CAI phase; with spinel and fassaite	1,2
	AOA forsterite	Defining major phase of AOAs	1,2
Garnet Group [$Ca_3(Al, Ti^{3+},V,Sc)_2Si_3O_{12}$]			
Rubinite ($Ca_3Ti^{3+}{}_2Si_3O_{12}$)	*CAI rubinite*	From type A CAIs	35
	URI rubinite	Significantly enriched in Y, Sc, and Zr; with eringaite	35
Eringaite ($Ca_3Sc_2Si_3O_{12}$)	*URI eringaite*	Occurs with rubinite	13,35
Melilite Group [gehlenite ($Ca_2Al_2SiO_7$) to åkermanite ($Ca_2MgSi_2O_7$)]			
	CAI melilite	Common in CAIs and Wark-Lovering rims	1,2

	AOA melilite	Gehlenite-rich, associated with anorthite and spinel	1,10
	URI melilite	Gehlenite-rich, associated with dacvisite, perovskite, spinel	26
Clinopyroxene Group [$Ca(Mg,Al,Ti^{3+},Sc,V)(Al,Ti^{4+},Si)SiO_6$			
Fassaite [$Ca(Mg,Al,Ti^{3+})(Si,Al,Ti^{4+})SiO_6$]	*CAI fassaite*	Common in CAIs and Wark-Lovering rims	1,2
	AOA fassaite	Common in association with forsterite and anorthite	10,11
	URI fassaite	Sub-millimeter grains with davisite, spinel, etc.	5,24,36
Burnettite ($CaV^{3+}AlSiO_6$)	*CAI burnettite*	Identified by from one fluffy type A CAI	37
Davisite ($CaSc^{3+}AlSiO_6$)	*URI davisite*	Associated with fassaite, perovskite, and spinel	12,15,24
Feldspar Group [$(Na,Ca)(Al,Si)_4O_8$]			
Anorthite ($CaAl_2Si_2O_8$)	*CAI anorthite*	Commonly as laths in CAIs; Primary anorthite is Na-free	1,2
	AOA anorthite	Common minor phase in AOAs	11
Dmisteinbergite ($CaAl_2Si_2O_8$)	*CAI dmisteinbergite*	A rare high-T, low-P polymorph of anorthite	38
Baghdadite [$Ca_3(Zr,Ti)Si_2O_9$]	*CAI baghdadite*	Known from a single 0.8-micron euhedral grain	39
Rhönite [$Ca_2(Mg,Al,Ti)_6(Si,Al)_6O_{20}$], Addibischoffite [$Ca_2(Al,Mg,V,Ti)_6(Al,Si)_6O_{20}$)], and Warkite [$Ca_2(Sc,Ti,Al,Mg,Zr)_6Al_6O_{20}$]			
	CAI rhönite	A rare constituent of A and B CAIs	40,41
	URI warkite	Micron-scale crystals with perovskite and davisite	26
Paqueite [$Ca_3TiSi_2(Al,Ti,Si)_3O_{14}$]	*CAI paqueite*	From a fluffy type A CAI	37
Thortveitite ($Sc_2Si_2O_7$)	*URI thortveitite*	Micron-scale crystals with fassaite and davisite	12

Silicate Glass (Ca,Mg,Al,Si,O)	*CAI silicate glass*	A component of hibonite-silicate spherules	2,42,43

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